Star Formation in our Extreme Galactic Center: Results from the CMZoom Survey

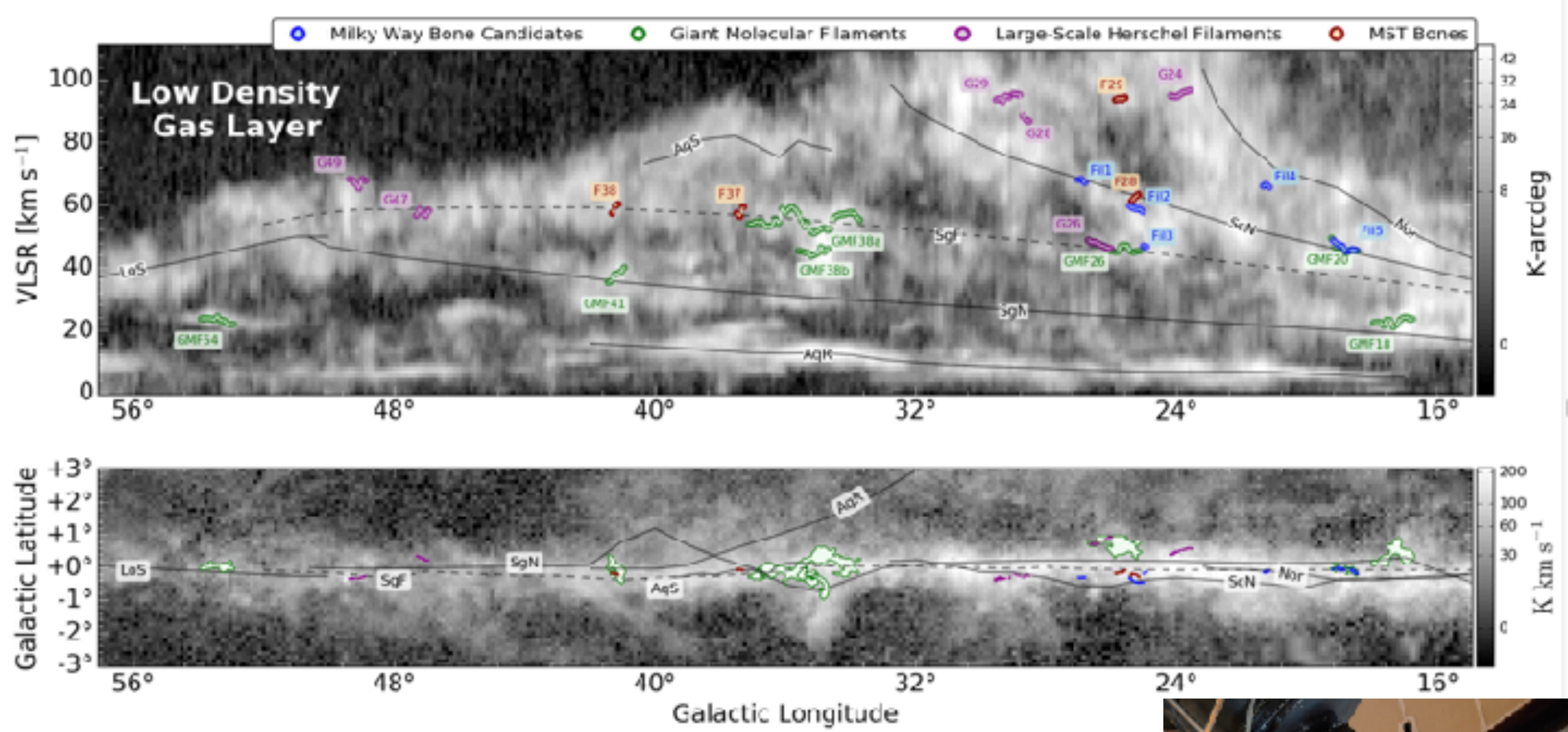
Cara Battersby

University of Connecticut

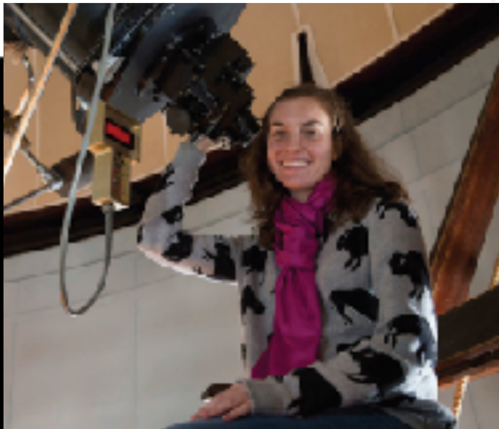
@battersbot

#TracingTheFlow

Updated full large Galactic filament analysis

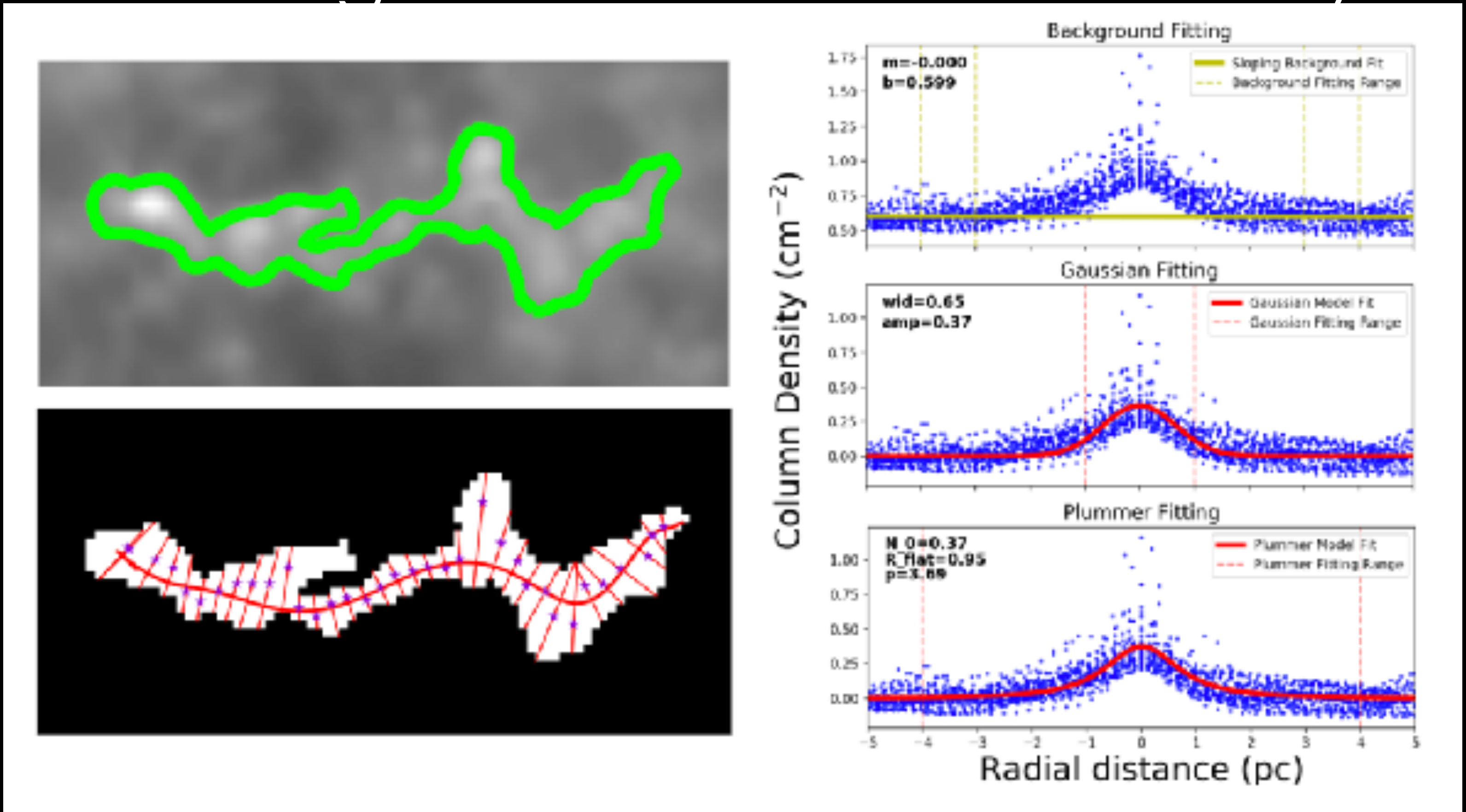


Catherine Zucker
Harvard PhD student



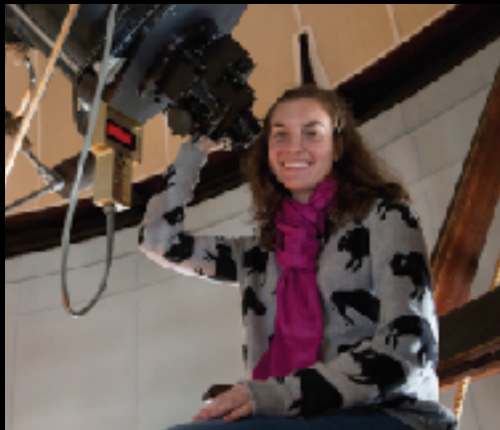
Zucker, Battersby, Goodman 2018

Updated full large Galactic filament analysis



Radfil! available now on GitHub

Catherine Zucker
Harvard PhD student



https://github.com/catherinezucker/radfil

Star Formation in our Extreme Galactic Center: Results from the CMZoom Survey

Cara Battersby

University of Connecticut

@battersbot

#TracingTheFlow

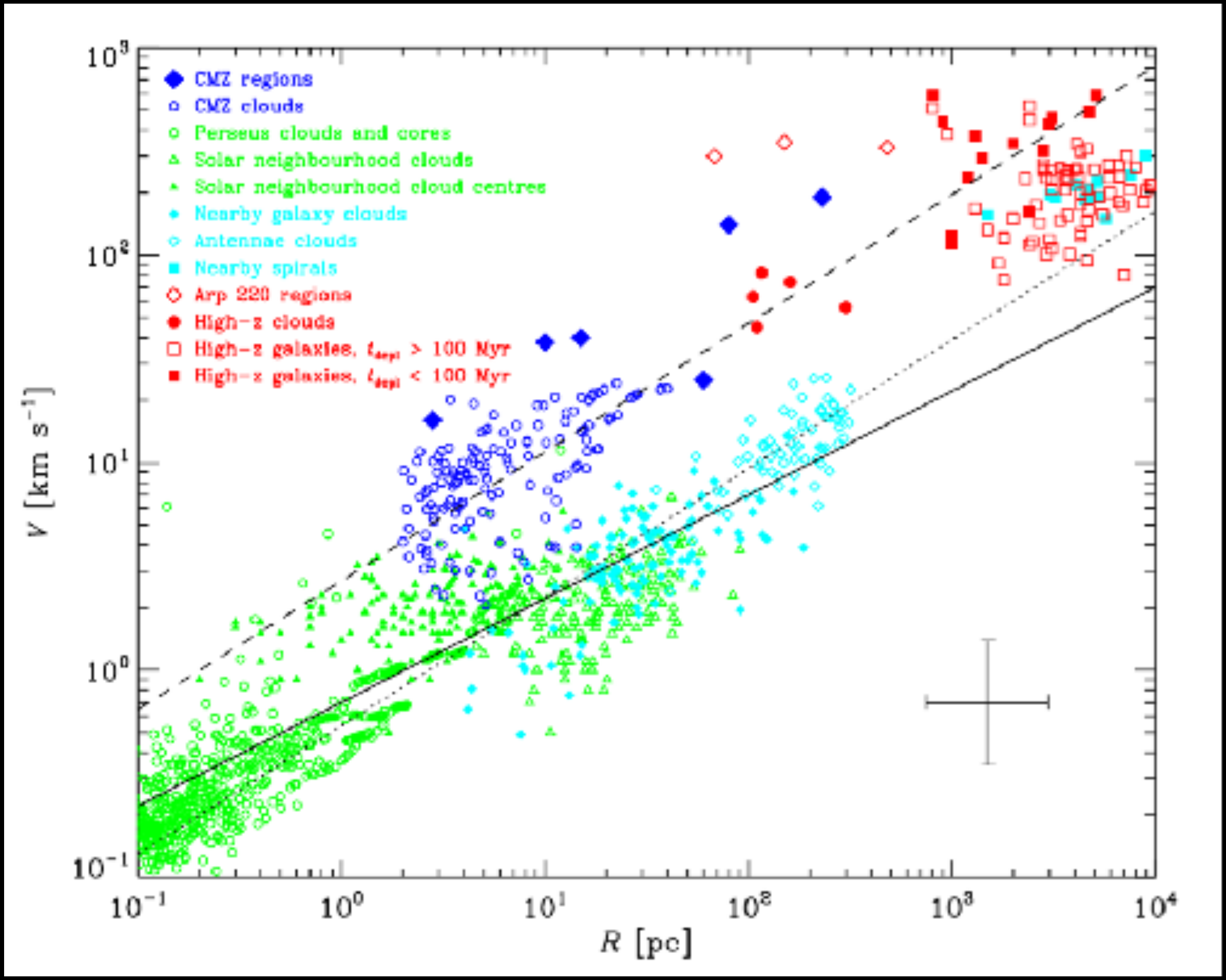
Central Molecular Zone

24 µm (Carey+ 2009.), 8 µm and 4.5 µm (Benjamin+20003)

Central Molecular Zone

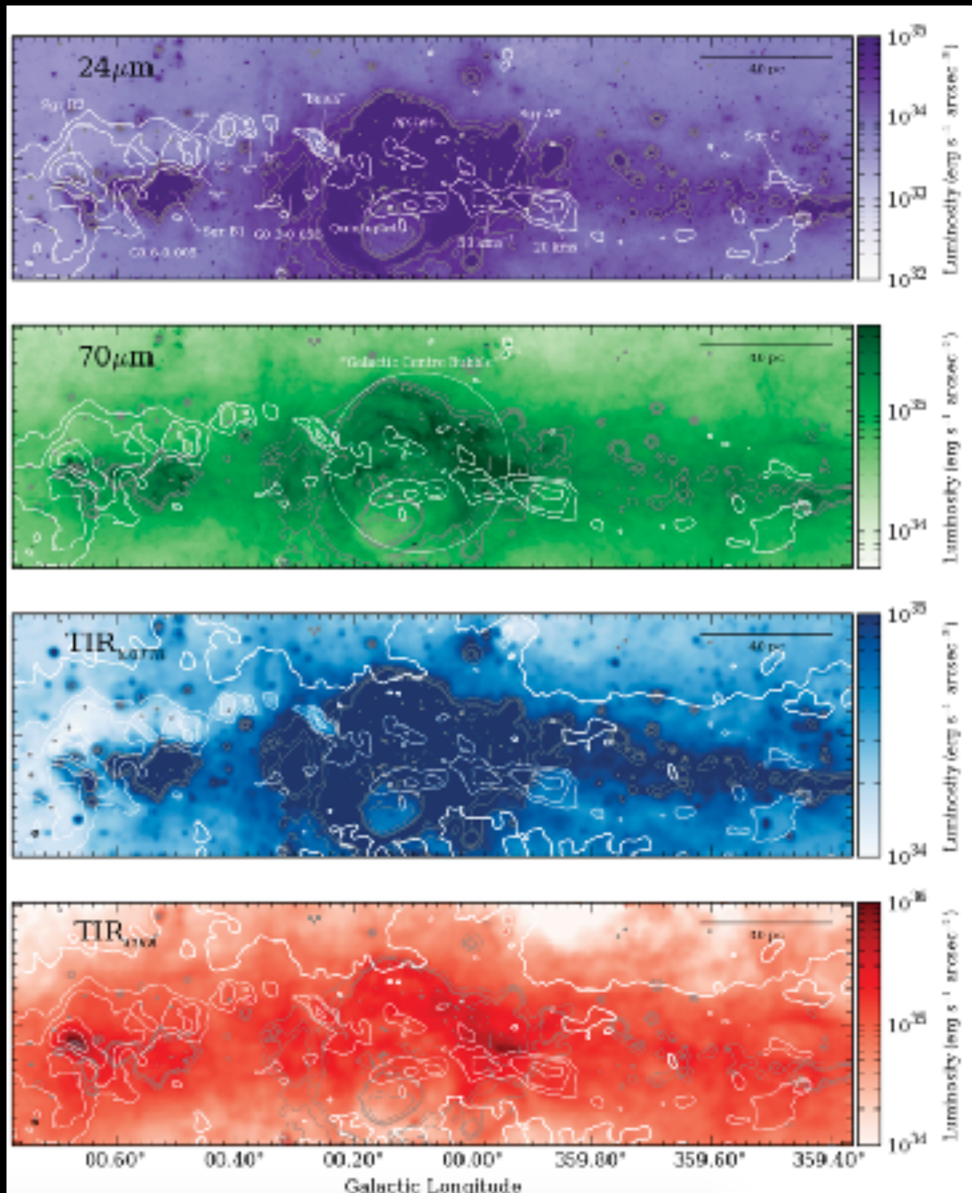
N(H_2) (Battersby+ in prep.), 70 µm (Hi-GAL, Molinari+2010,2011), 8 µm (GLIMPSE, Benjamin+20003)

The Central Molecular Zone: A window into the distant universe



Kruijssen & Longmore 2013

Star Formation Rates of the CMZ



Ash Barnes
Liverpool PhD! ~~student~~
ITA Postdoc

Barnes et al. 2017

Star Formation Rates of the CMZ

Tracer
24 μm IR luminosity
70 μm IR luminosity
Total IR luminosity
YSO counting
free-free emission

About **0.06 - 0.10 M⊙/year**

for a wide variety of methods and tracers

— not underestimating SFR based on one method

[a] Approximately |ℓ|<1° and |b|<0.5°

[b] Contaminated by main-sequence stars (see Koepferl+2015)

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Barnes et al. 2017

Star Formation Rates of the CMZ

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24 μm IR lumino
70 μm IR lumino
Total IR luminosi
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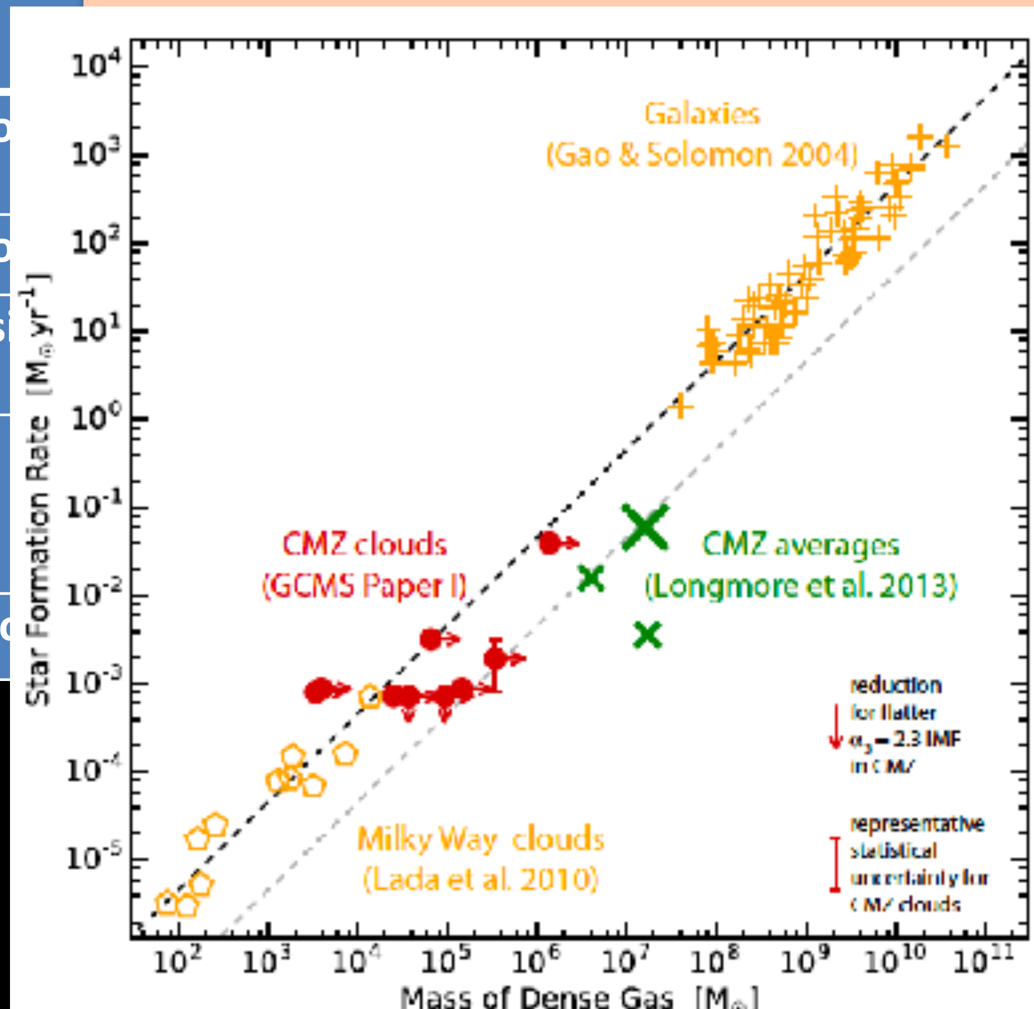
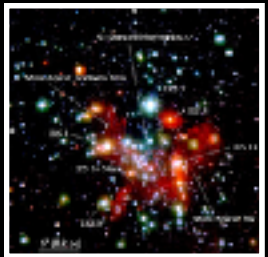
and

ased

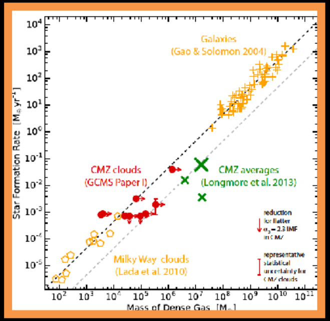


Kauffmann et al. 2016

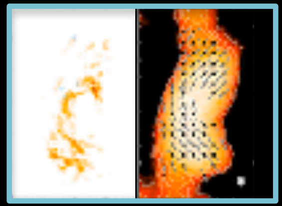
Star Formation in the CMZ

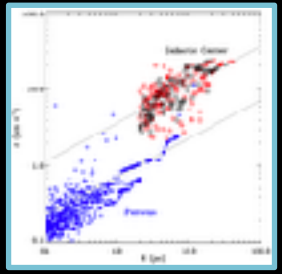


There are many extreme stars and clusters in the CMZ

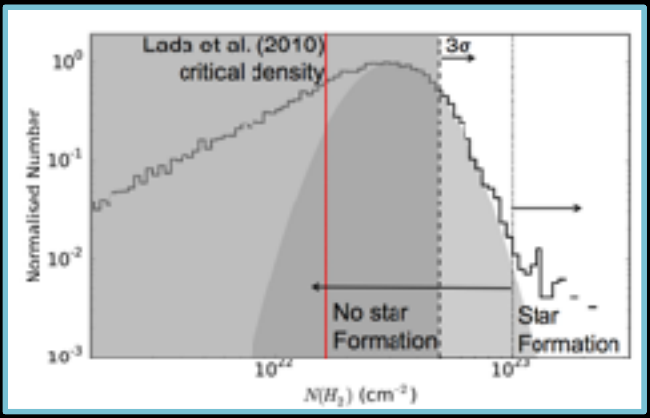


The CMZ is currently underproducing stars by ~10





CMZ gas is hot, dense, chemically complex, turbulent, with strong B fields, and the ISRF and CRIR are high → any of these may affect SF



There is NO universal density threshold for Star Formation — but maybe an environmentally dependent one

***Figures**: ESO/VLT of Young Nuclear Cluster, Brick: Rathborne et al. 2014, Pillai et al. 2015, Dense gas relation: Kauffmann et al. 2016, size-linewidth: Shetty et al. 2012, CMZoom SF threshold: CMZoom in prep.*

SMA Legacy Survey of the Central Molecular Zone

- Large primary beam + wide bandwidth + long wavelength + high angular resolution → detect early star formation across a large area
- First survey of the CMZ ever to be able to do so

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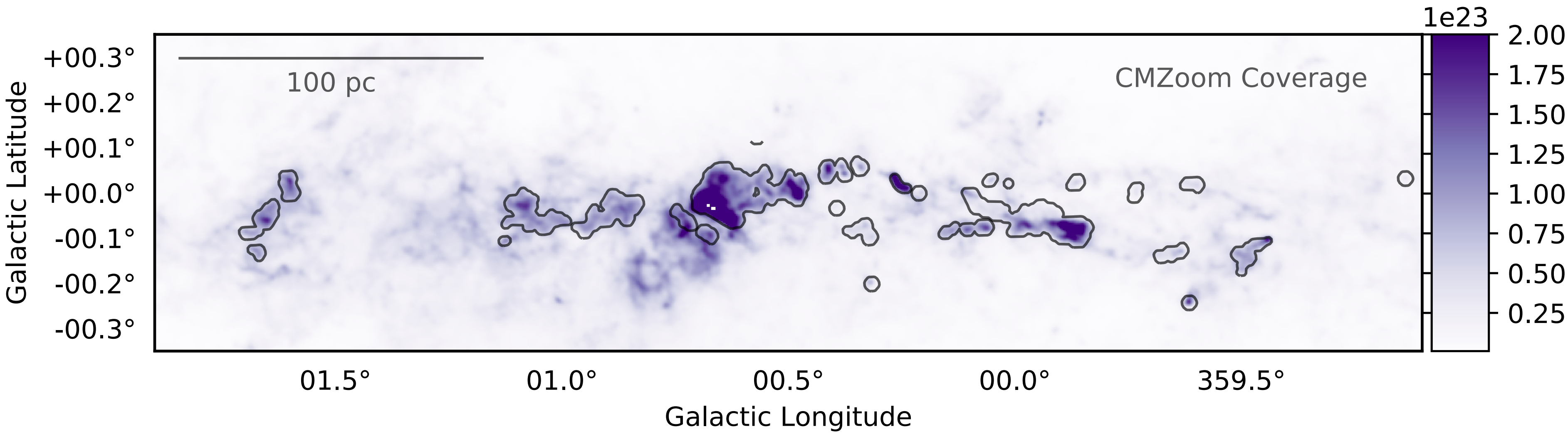
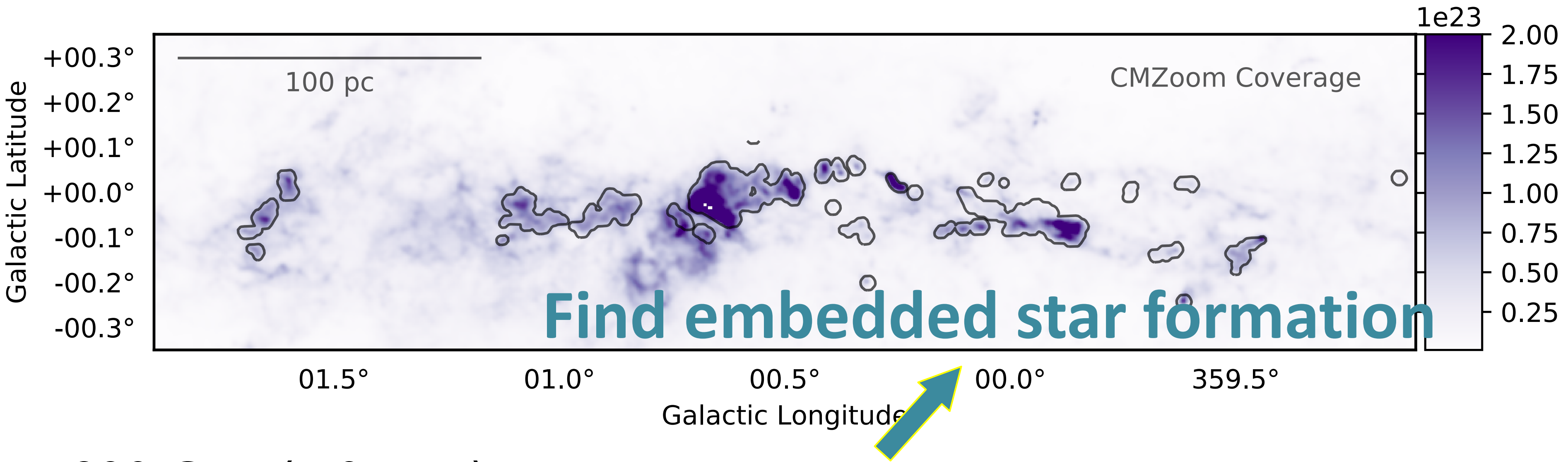
CMZoom



- 230 GHz (1.3 mm)
- 240 arcmin^2 (above N(H_2) = 10^{23} cm^{-2} or 3x10^{22} cm^{-2})
- 4'' (0.2 pc) resolution, Δv~1.1 km/s
- dust continuum + spectral lines (H_2CO, ^{12}CO, ^{13}CO, $C^{18}O$, SiO, CH_3OH, CH_3CN, etc.): 8+ GHz bandwidth
- 3 mJy RMS continuum, 0.4 K
- 550 hours (50 subcompact, 450 compact/custom) over 4 yrs
- Complement with single-dish (APEX, CSO) observations

CMZoom



- ⋆ 230 GHz (1.3 mm)
- ⋆ 240 arcmin^2 (above N(H_2) = 10^{23} cm^{-2} or 3×10^{22} cm^{-2})
- ⋆ 4'' (0.2 pc) resolution, $\Delta v \sim 1.1$ km/s
- ⋆ dust continuum + spectral lines (H_2CO, ^{12}CO, ^{13}CO, $C^{18}O$, SiO, CH_3OH, CH_3CN, etc.): 8+ GHz bandwidth
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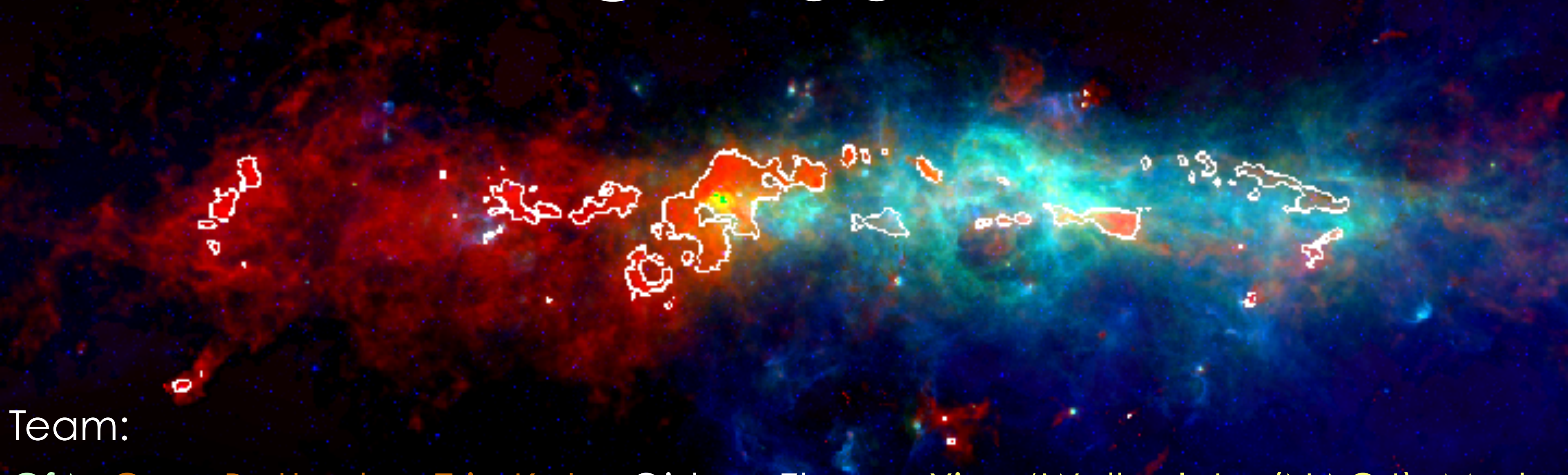
CMZoom



Team:

CfA: Cara Battersby, Eric Keto, Qizhou Zhang, Xing 'Walker' Lu (NAOJ), Mark Graham (Oxford), Nimesh Patel, Volker Tolls, Dennis Lee, Jimmy Castaño, Liz Gehret, Irene Vargas-Salzar, Perry Hatchfield, Daniel Callanan, Elizabeth Gutierrez

Bonn: Jens Kauffmann, Thushara Pillai

Liverpool: Steve Longmore, Daniel Walker (CfA), Jonny Henshaw

University of Colorado, Boulder: John Bally

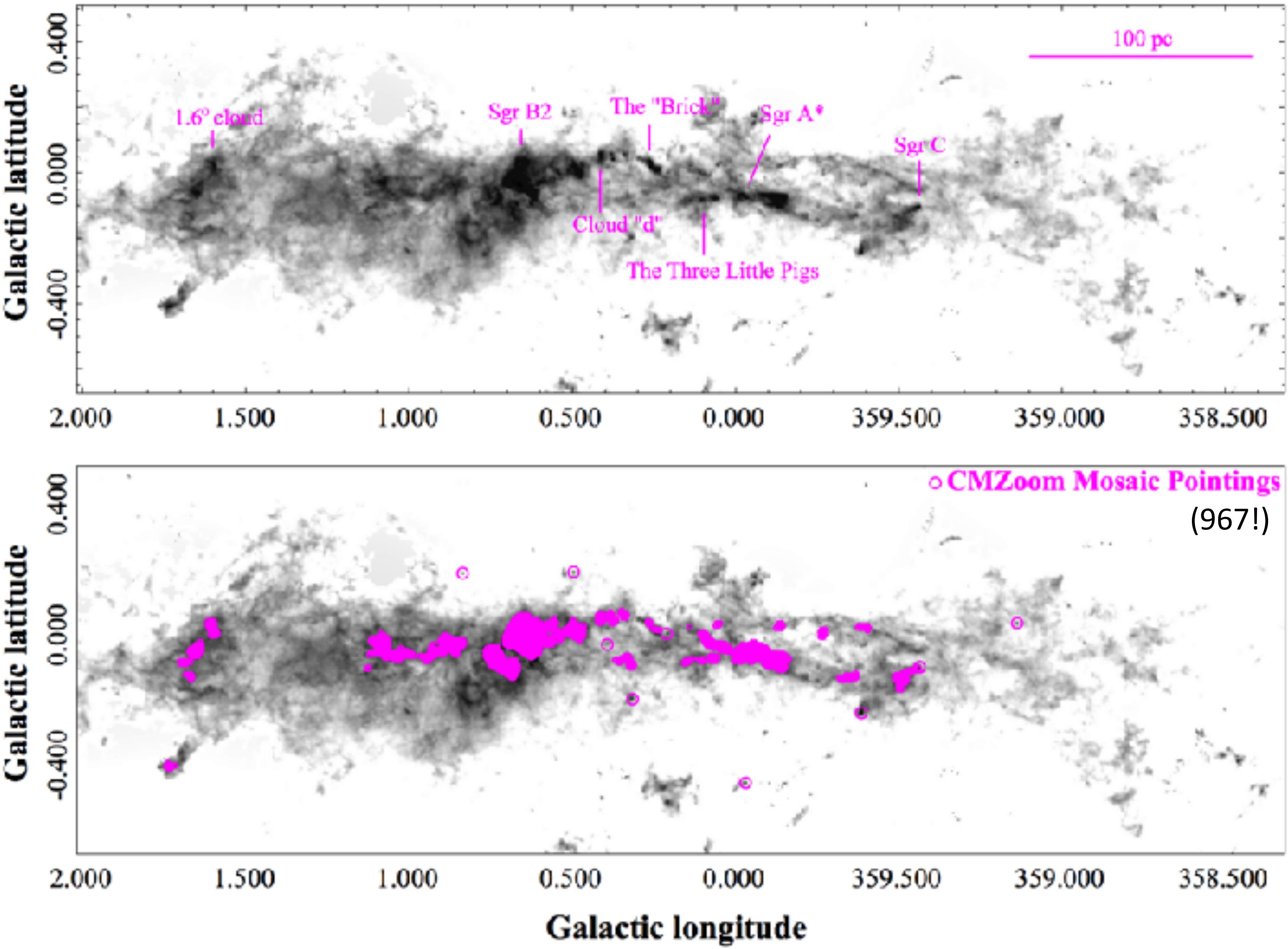
Heidelberg: Diederik Kruijssen, Ash Barnes NRAO: Betsy Mills, Natalie Butterfield ESO: Adam Ginsburg, Katharina Immer, Leeds: Katharine Johnston

Peking University: Luis C. Ho, Perth: Andrew Walsh

CMZoom

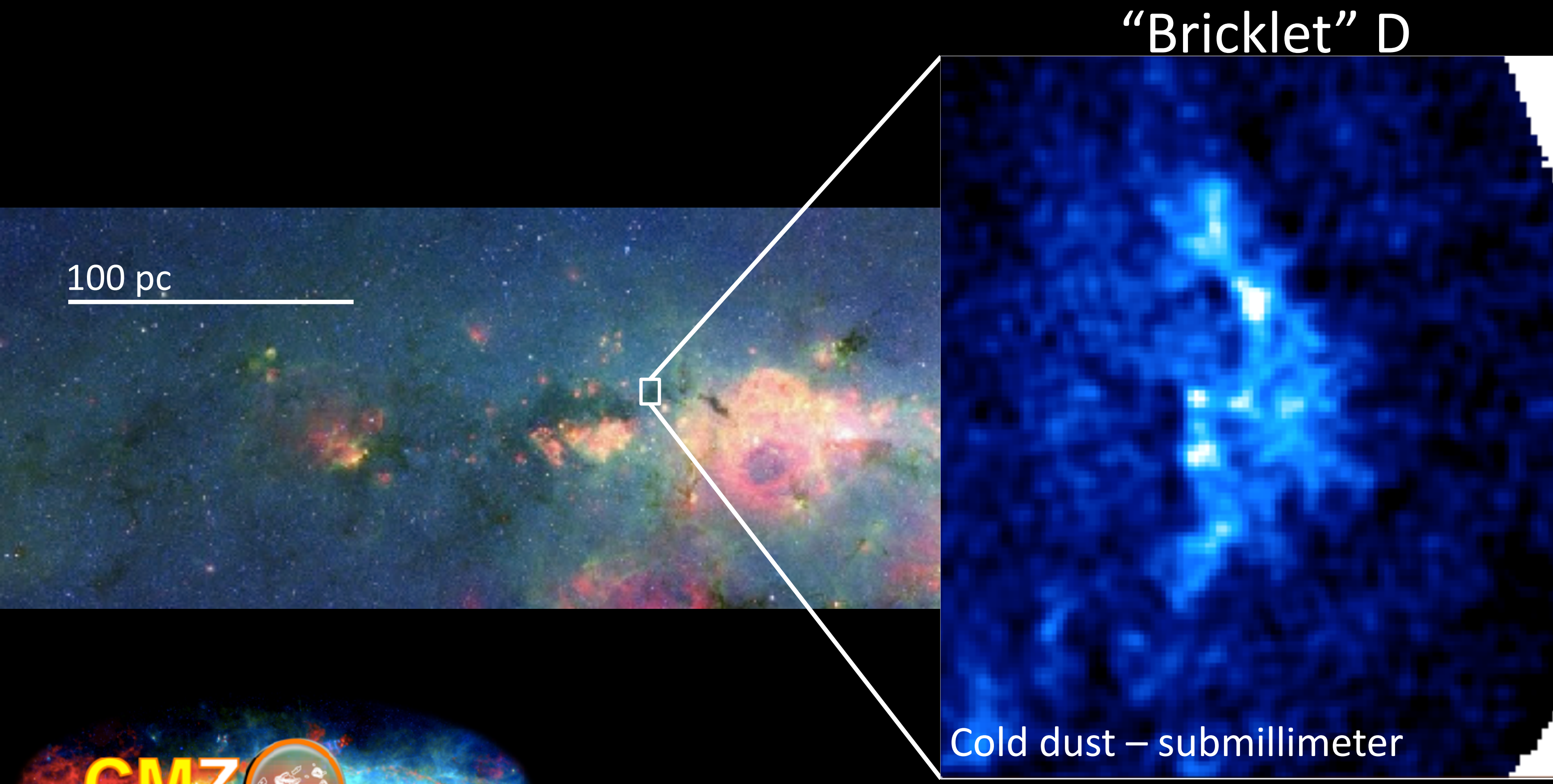
Team

CMZoom



100 pc
1.6° cloud
Sgr B2
The "Brick"
Sgr A*
Sgr C
Cloud "d"
The Three Little Pigs
Galactic latitude
0.400
0.000
-0.400
2.000
1.500
1.000
0.500
0.000
359.500
359.000
358.500
CMZoom Mosaic Pointings
(967!)
Galactic longitude

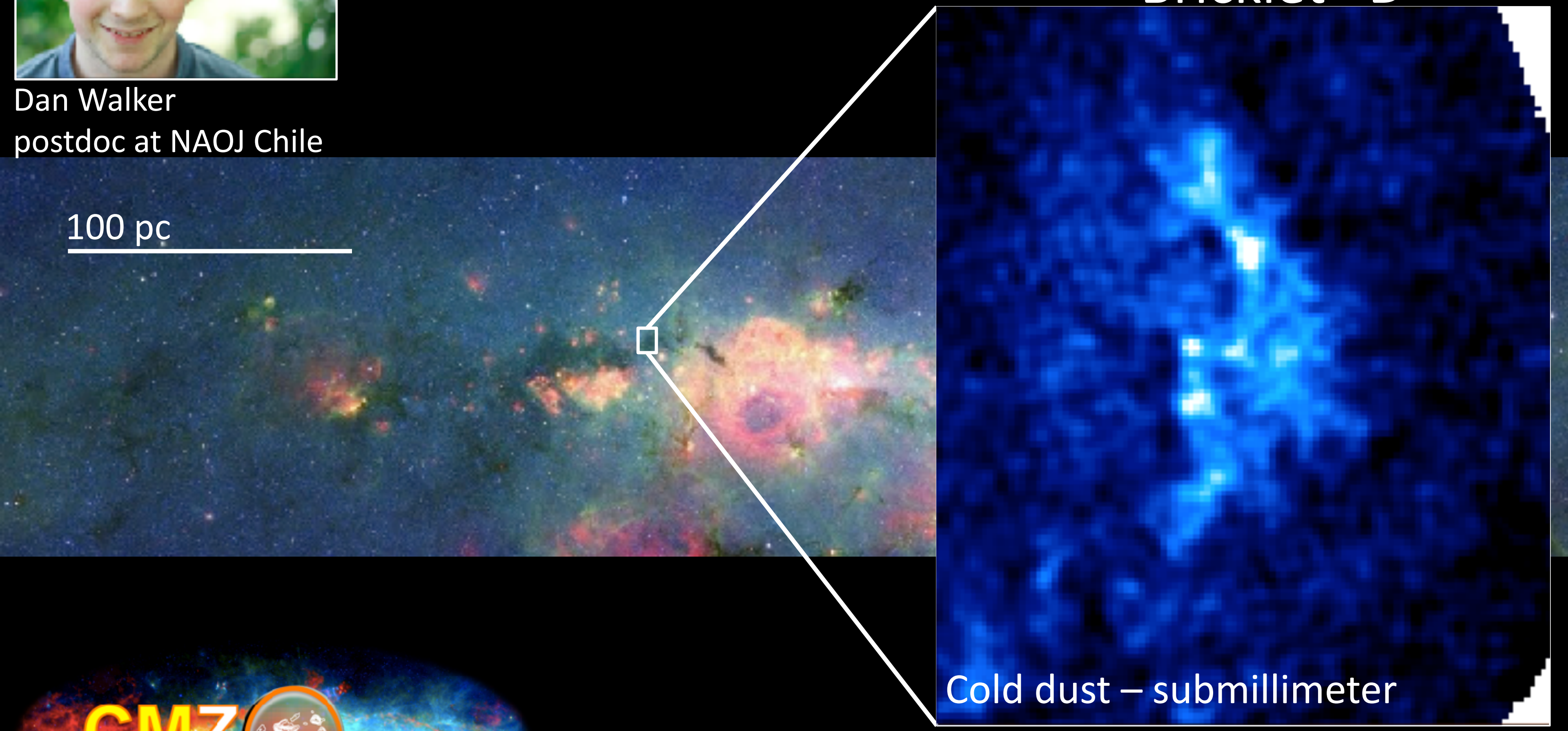
Central Molecular Zone



Central Molecular Zone

Dan Walker
postdoc at NAOJ Chile

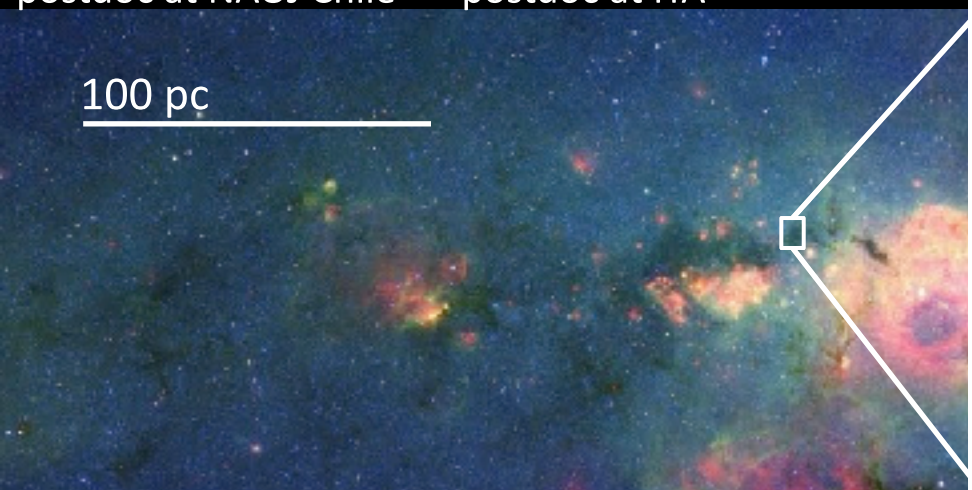
"Bricklet" D

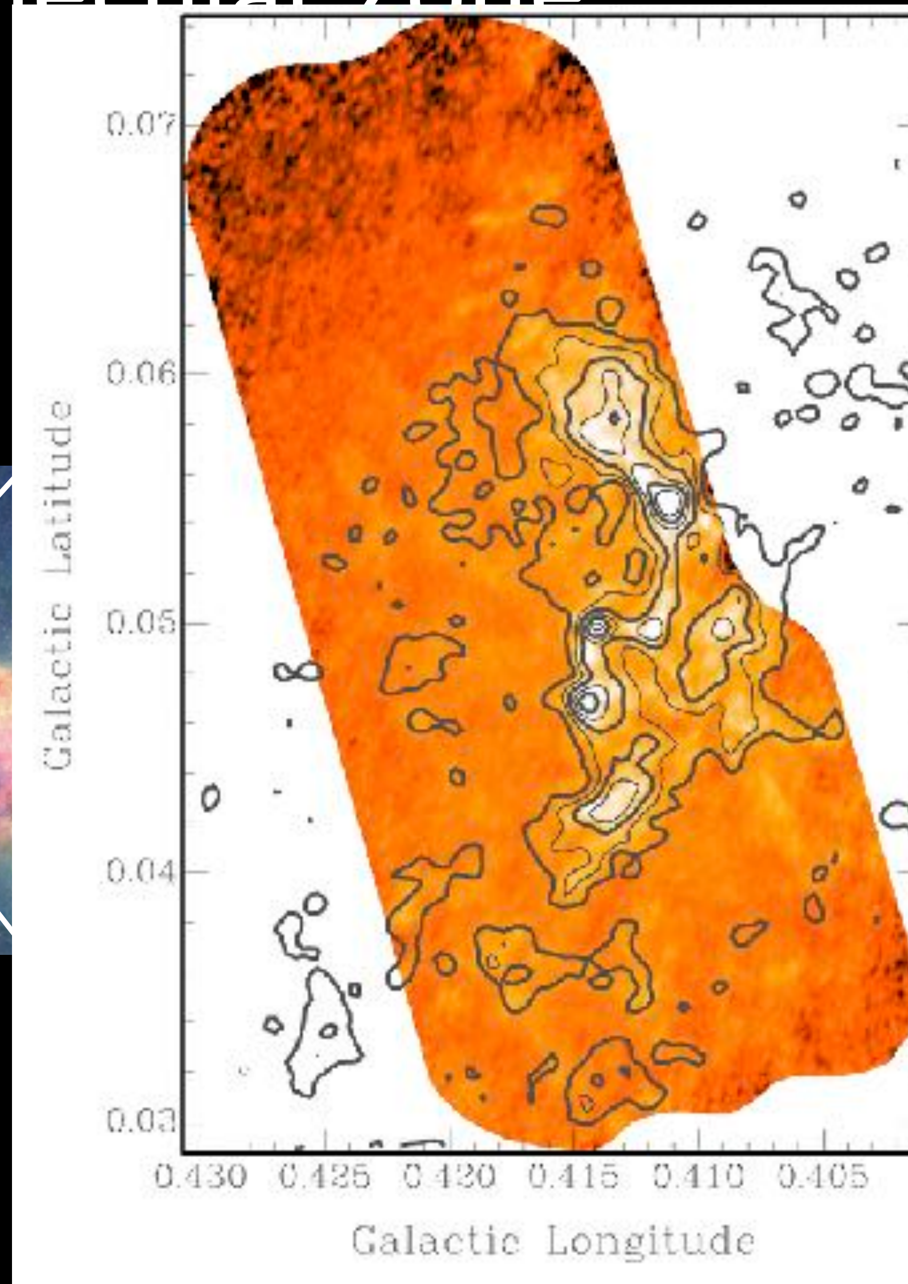


Central Molecular Zone

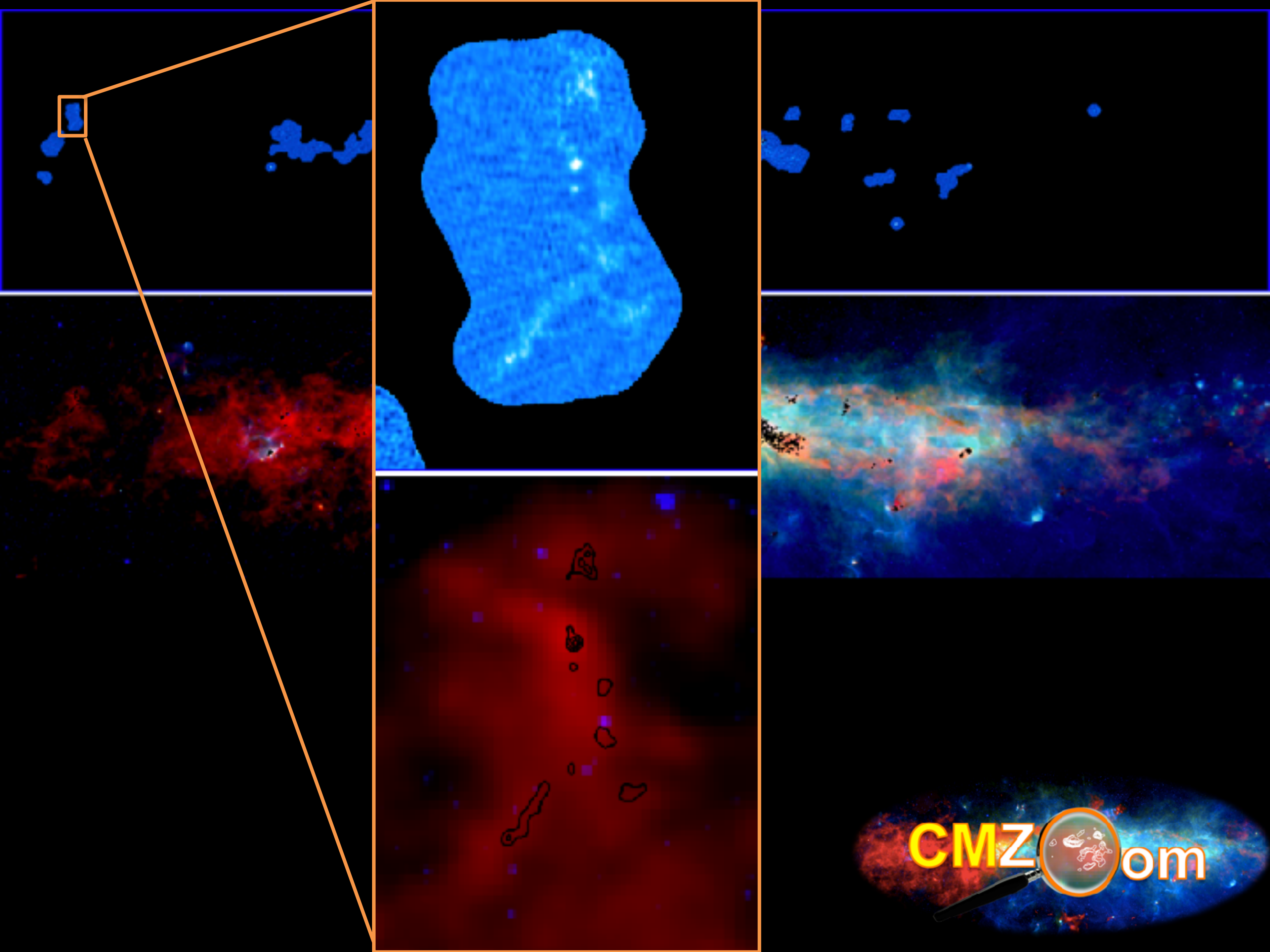
Dan Walker
postdoc at NAOJ Chile

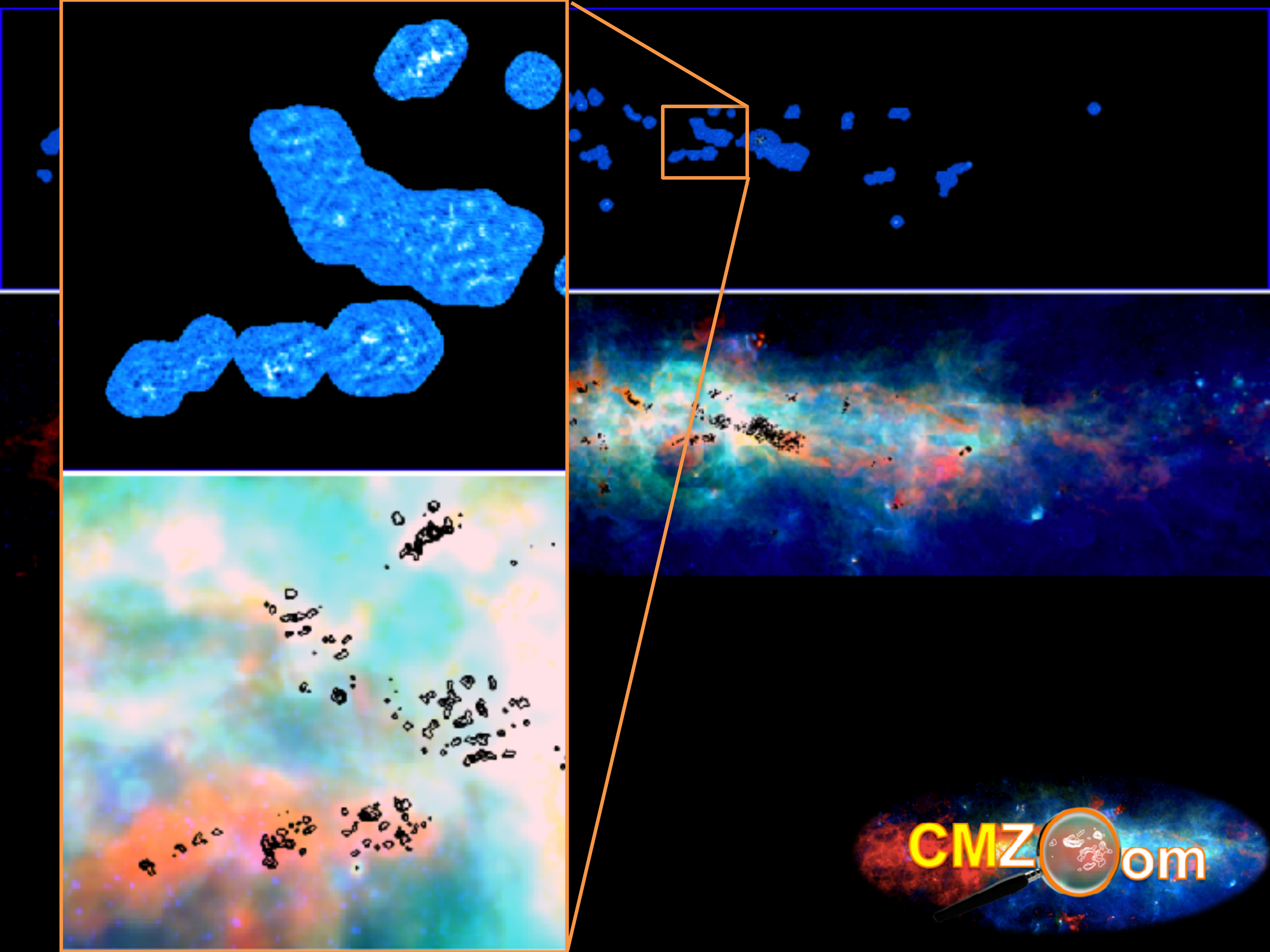
Ash Barnes
postdoc at ITA





$N(H_2)$ from HiGAL Battersby+, in prep., 70 μm from HiGAL, Molinari+ 2011, 8 μm from GLIMPSE (Benjamin+ 2003)


CMZoom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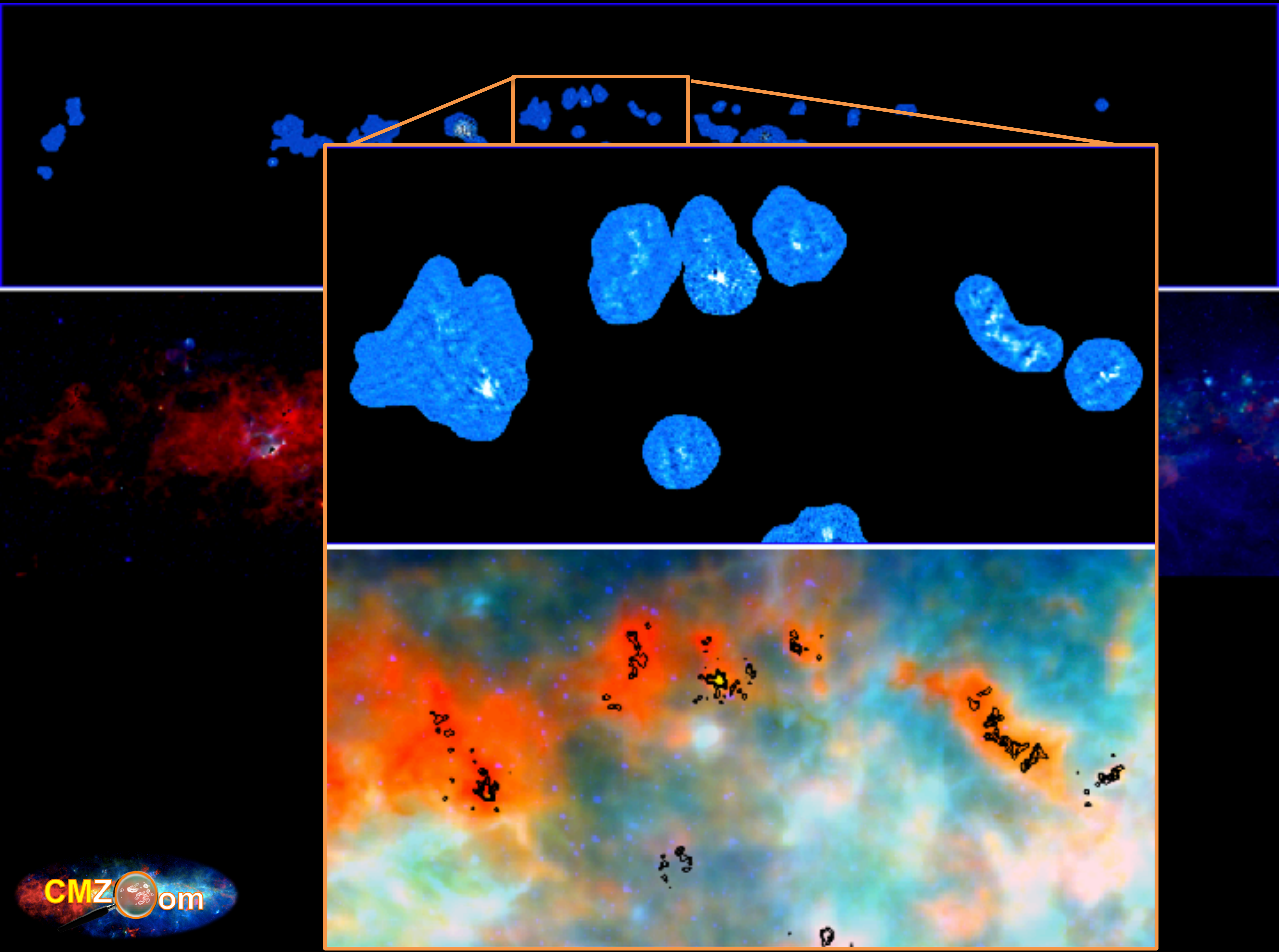
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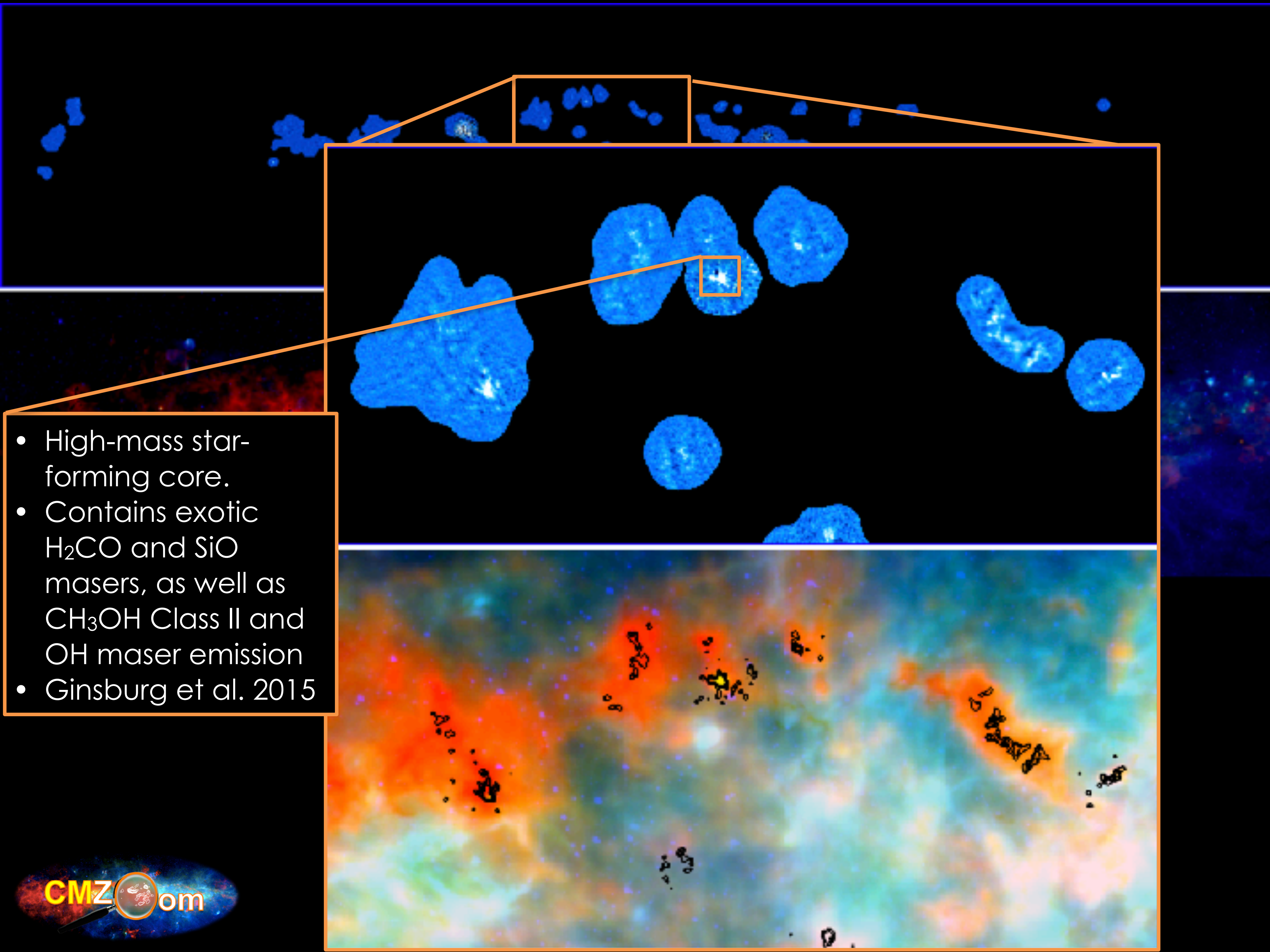
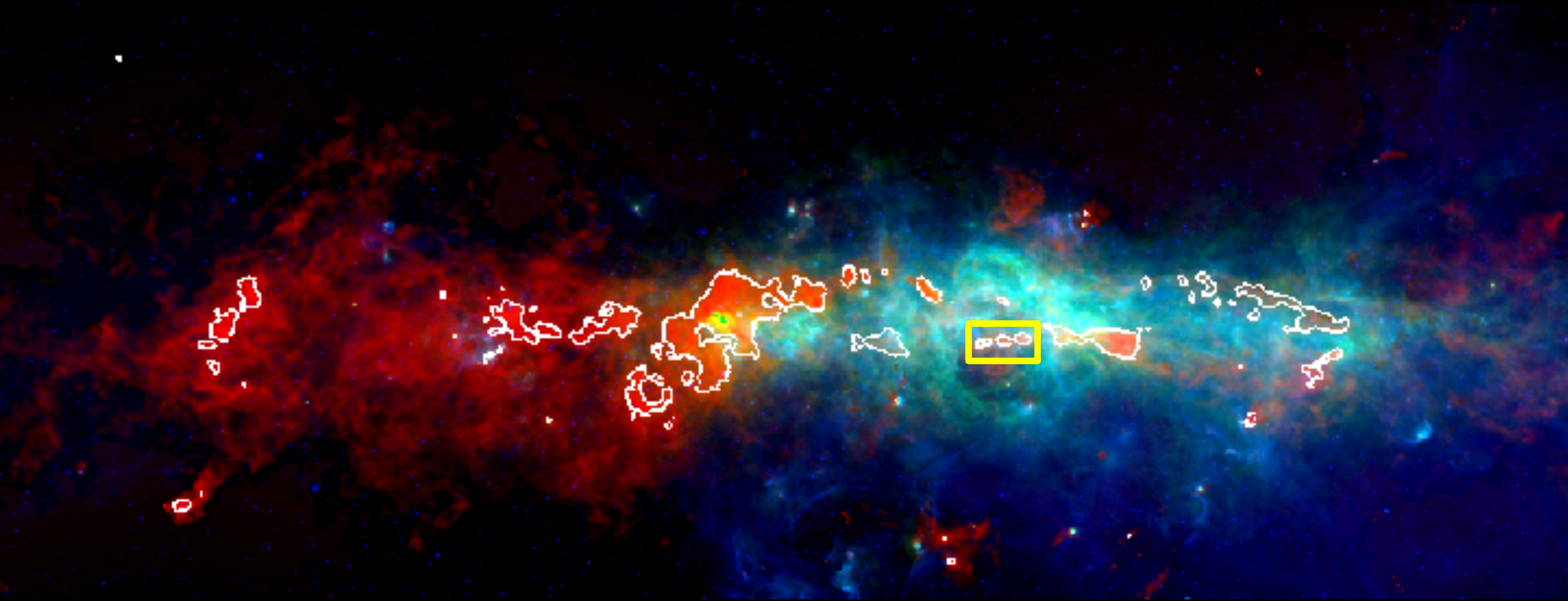
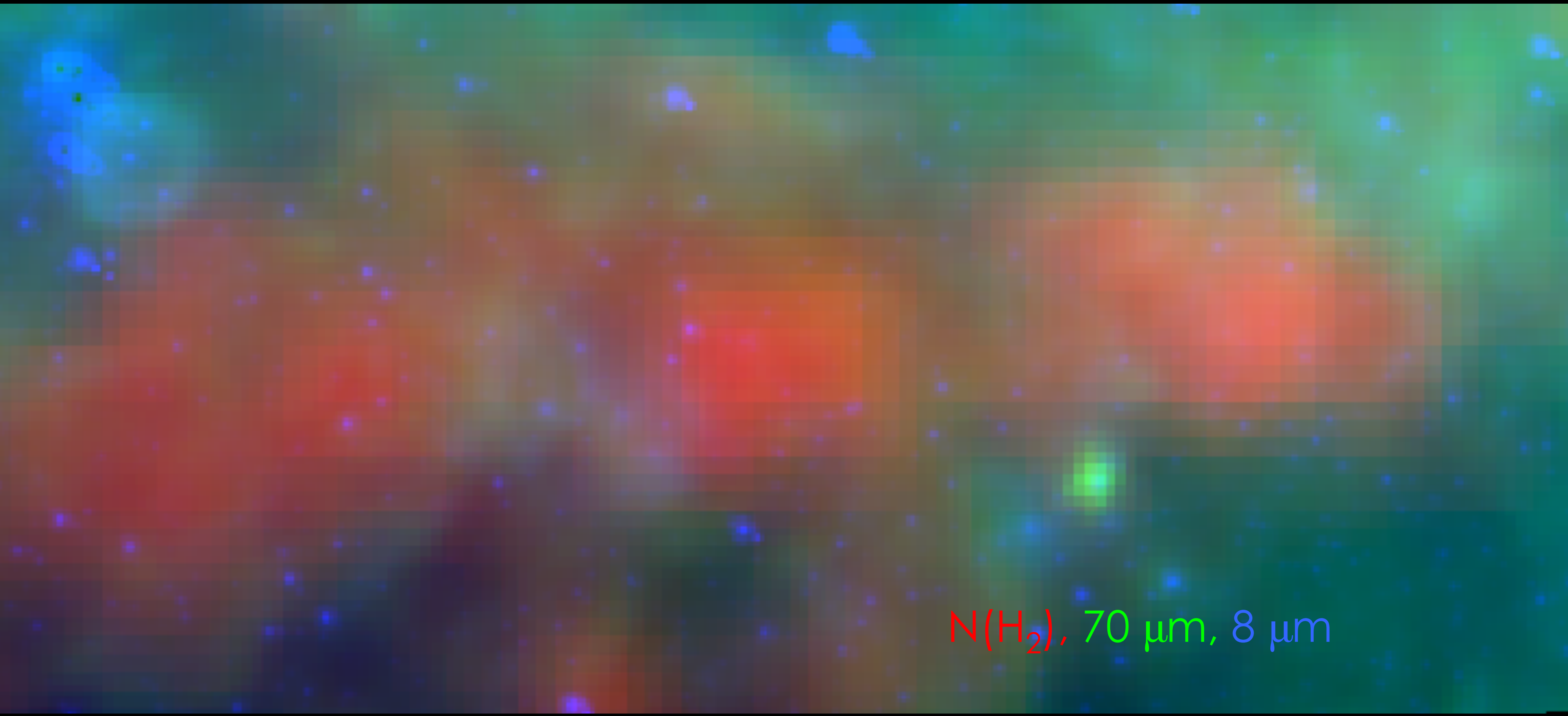


- High-mass star-forming core.
- Contains exotic H_2CO and SiO masers, as well as CH_3OH Class II and OH maser emission
- Ginsburg et al. 2015

Star Formation in the CMZ



Why is the SFR low in the CMZ?



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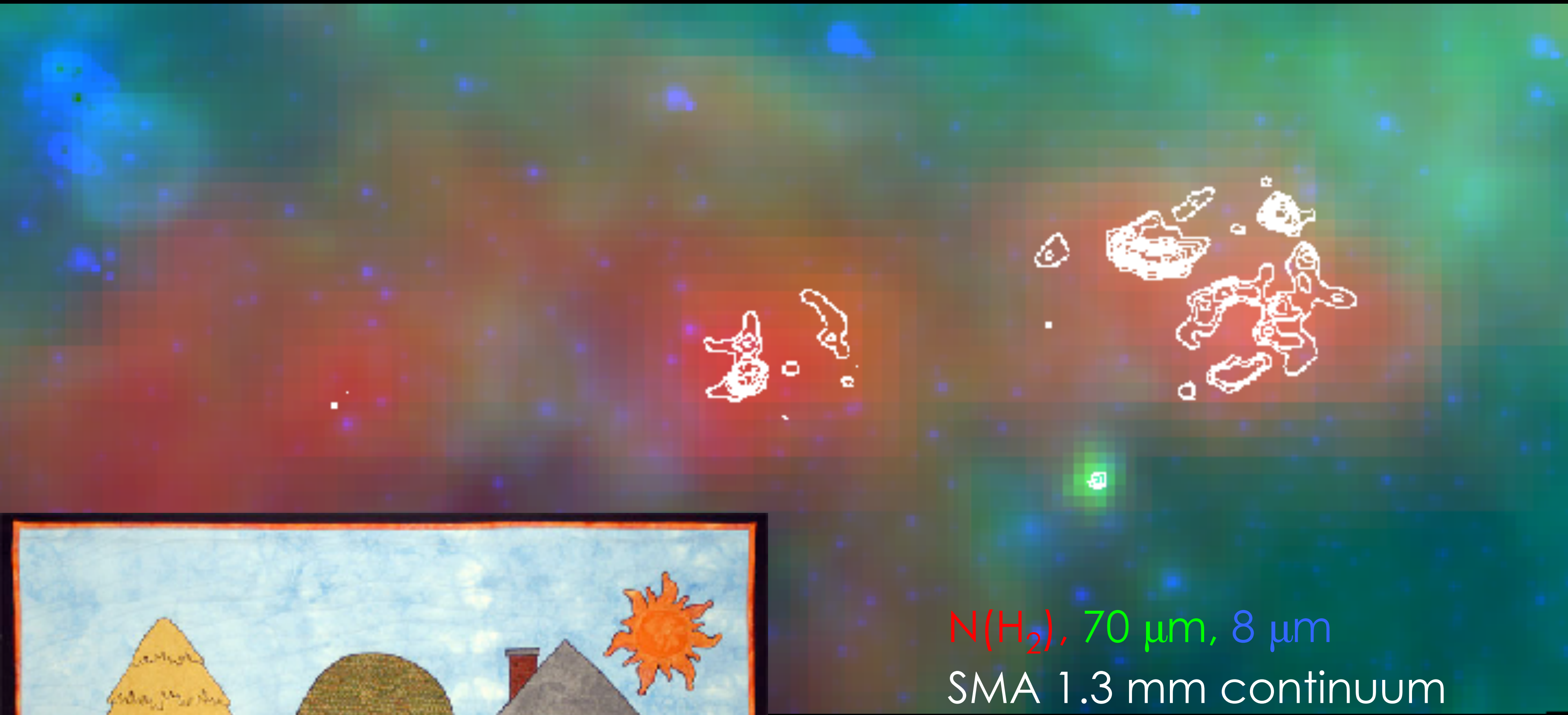

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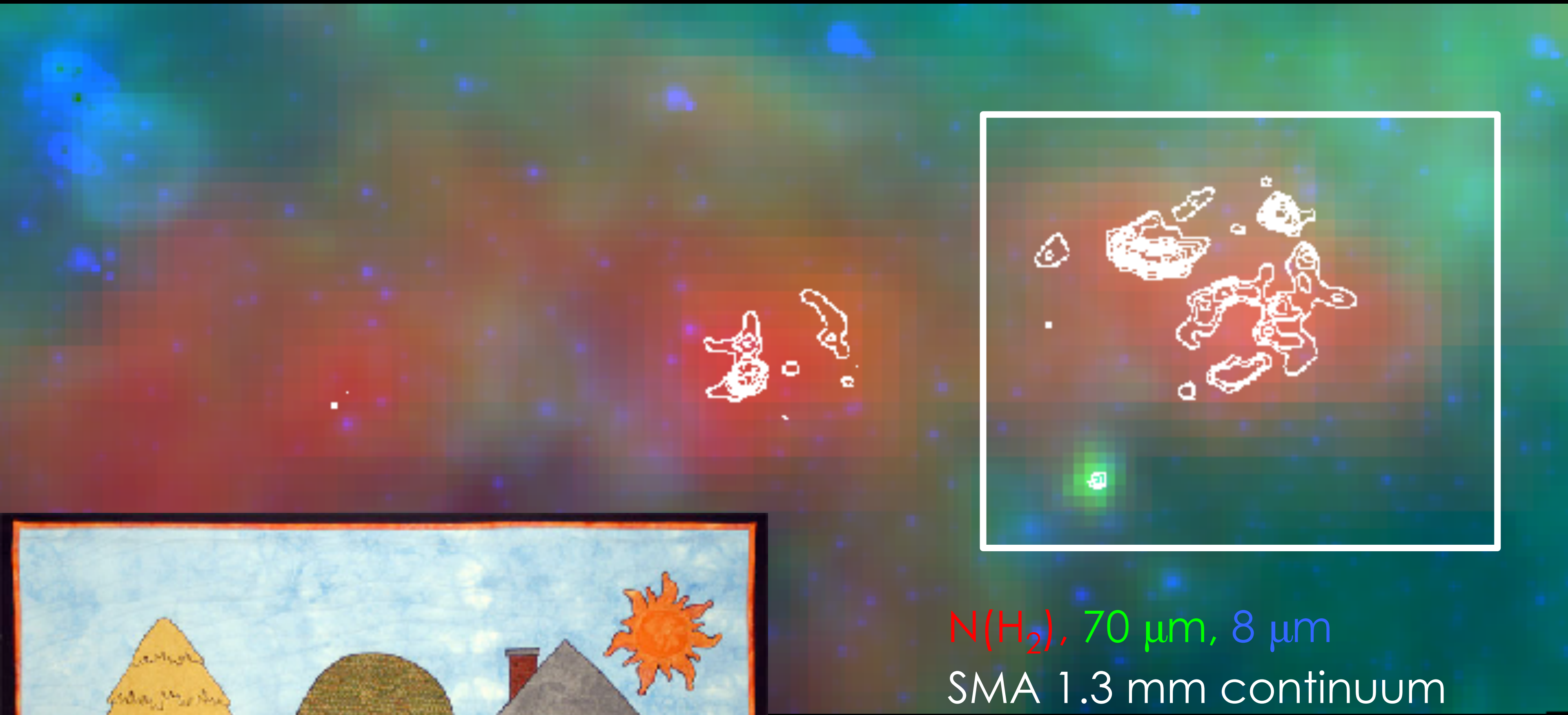

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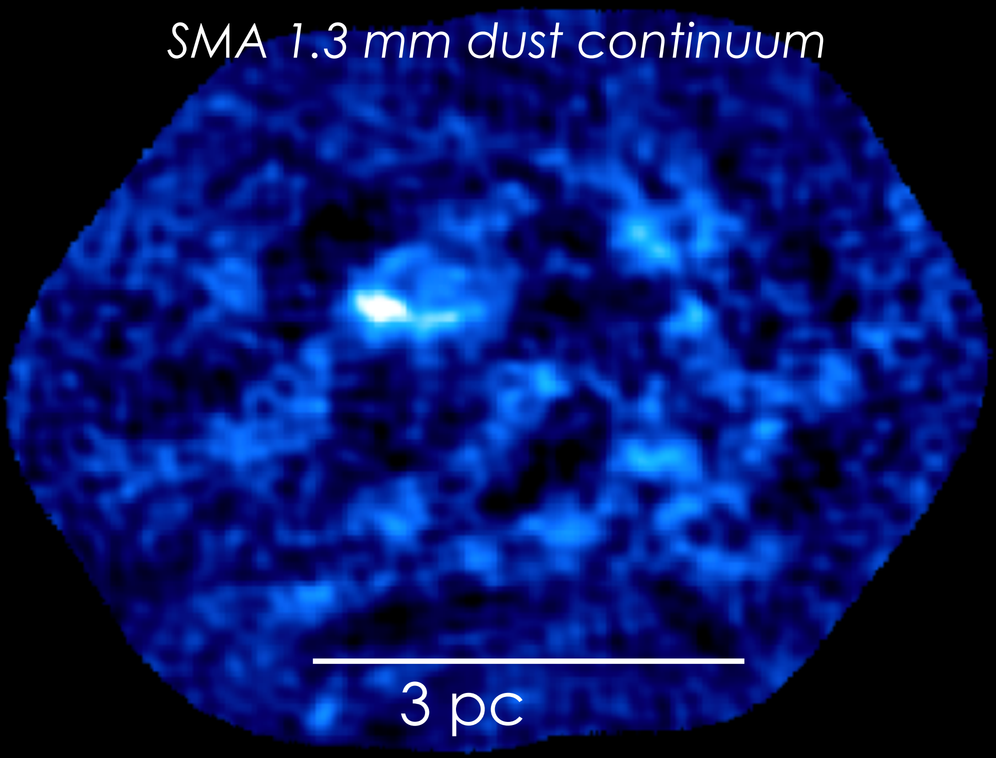

Is it star forming?

- ✓ Dense gas
- ✓ Shocked, highly excited gas
- ❑ Virial ratio < 2
- ❑ Power-law tail in N-PDF
- ❑ Outflow, localized hot-core chemistry, masers, UCHII regions…

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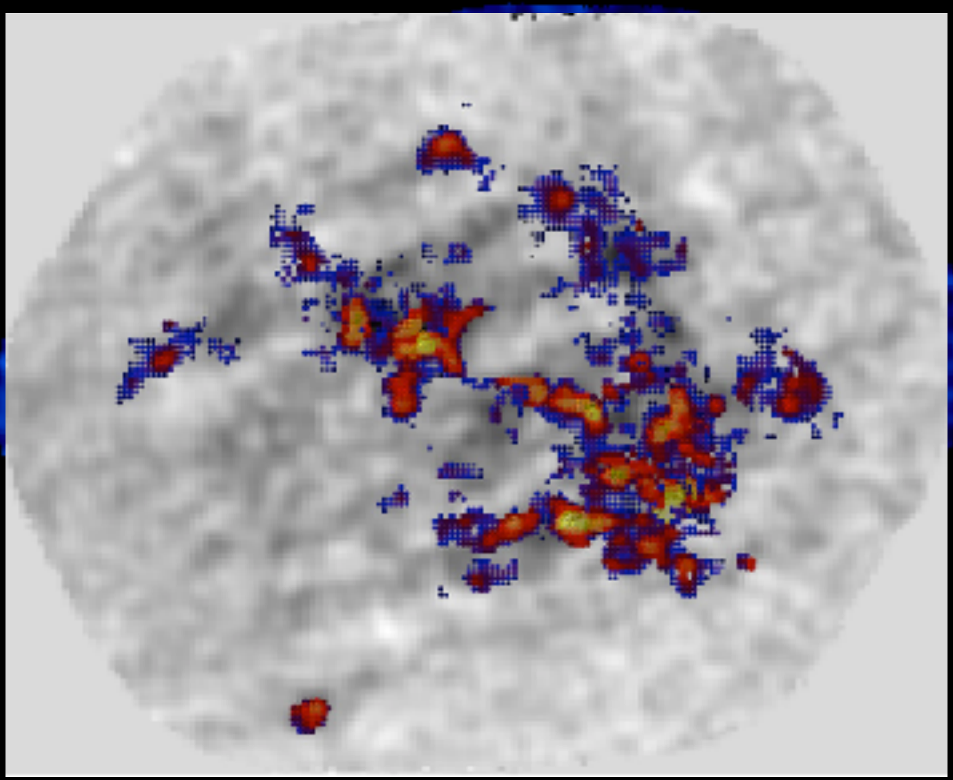
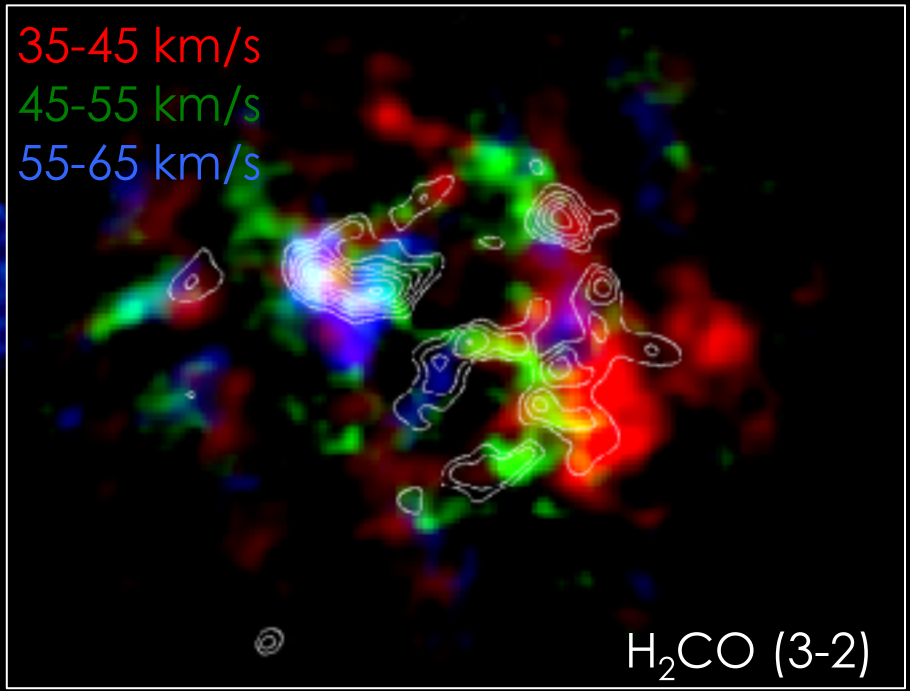
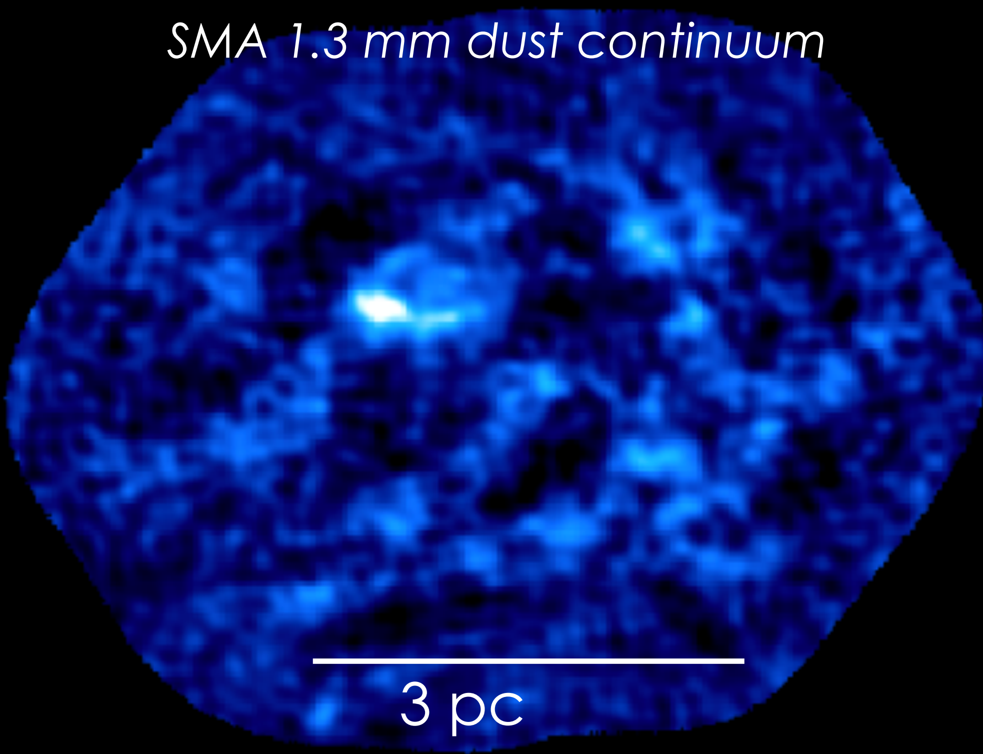




SCOUSE line fitting

Jonny Henshaw, MPIA

Why is the SFR low in the CMZ?



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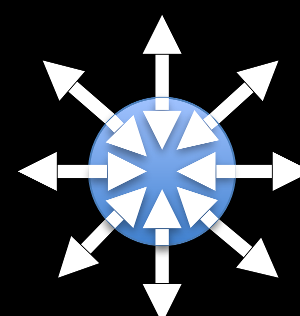
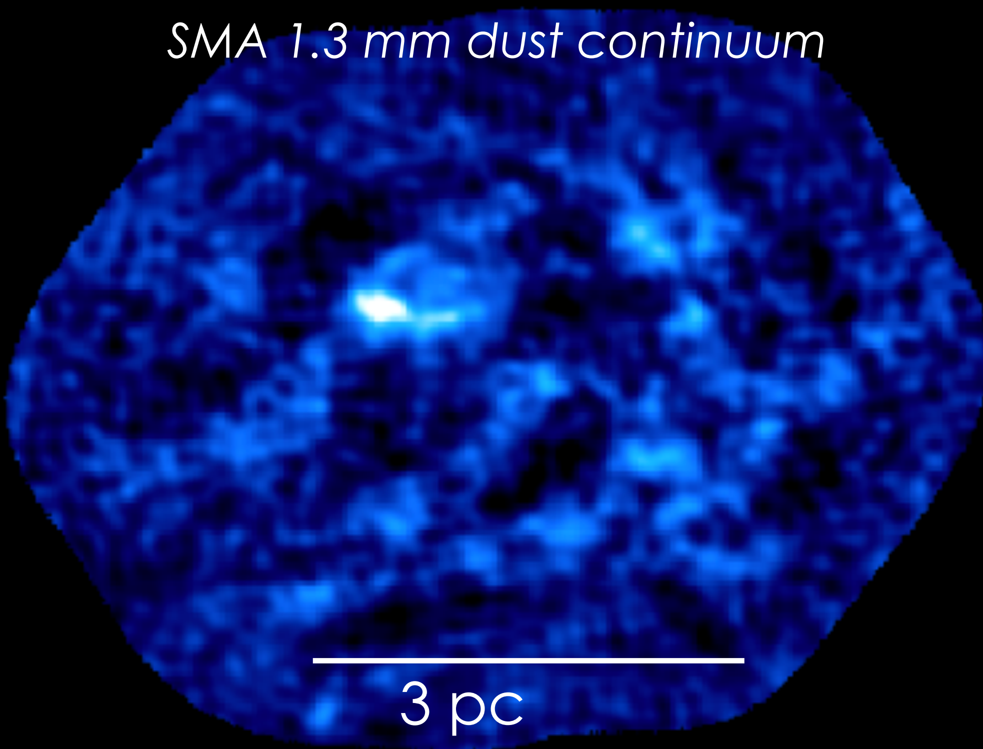


Gravity vs. pressure (thermal and turbulence)

CMZoom

Why is the SFR low in the CMZ?



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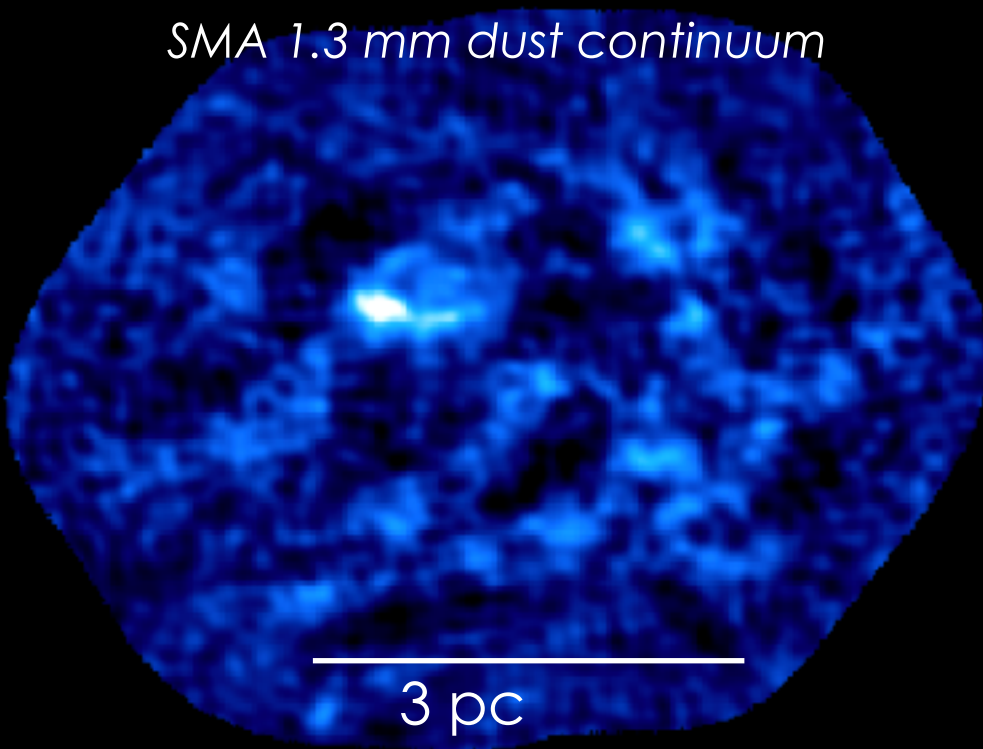
CMZoom

Why is the SFR low in the CMZ?

SMA 1.3 mm dust continuum



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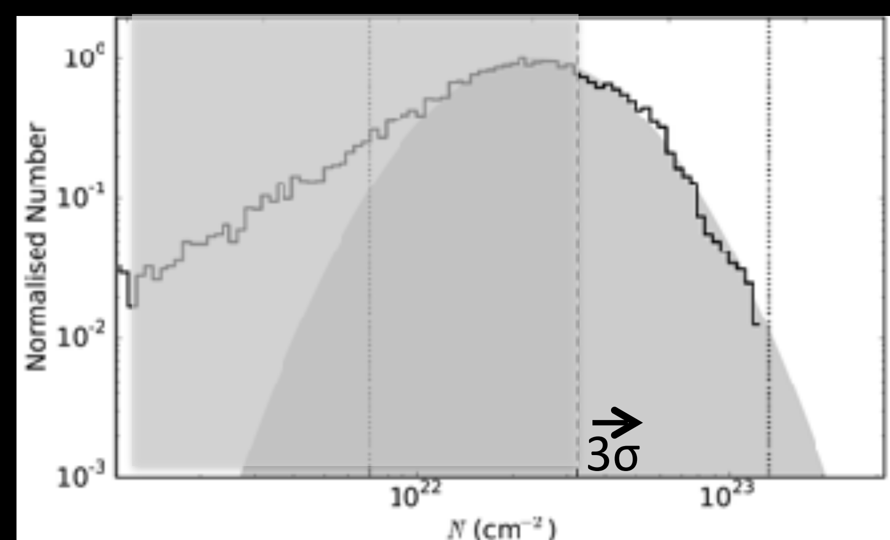

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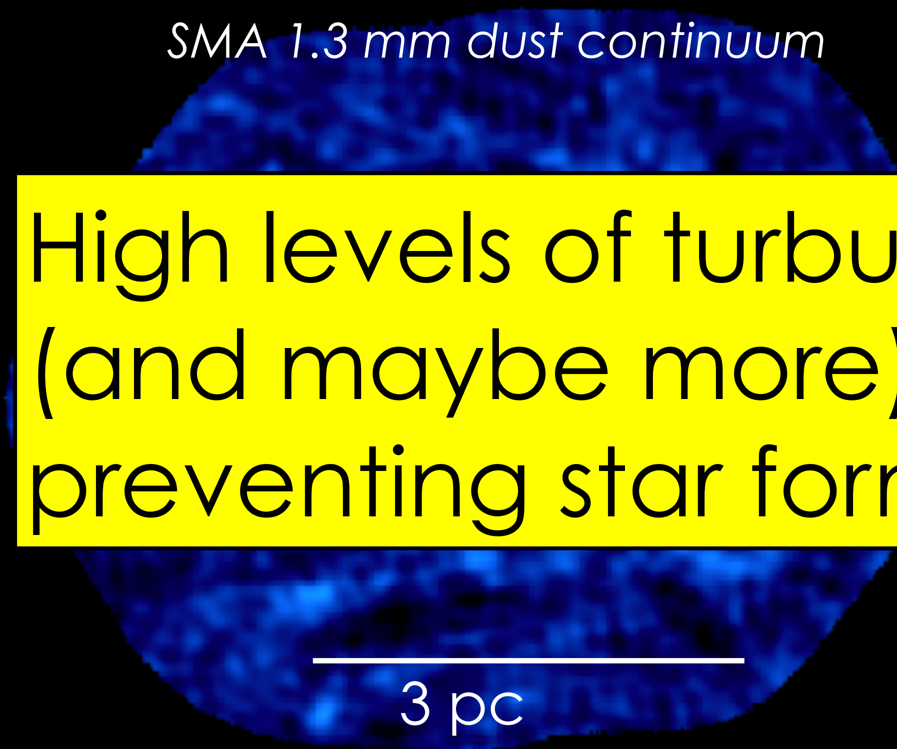

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✓ Dense gas

✓ Shocked, highly excited gas

☒ Virial ratio < 2

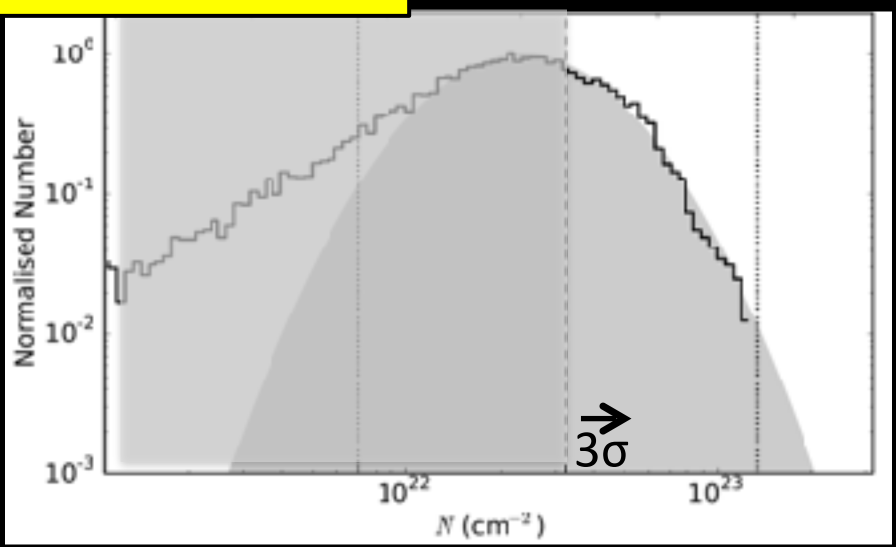
…w tail in N-PDF

… localized hot-

…emistry, masers,

…gions…

High levels of turbulence[1] (and maybe more) are preventing star formation

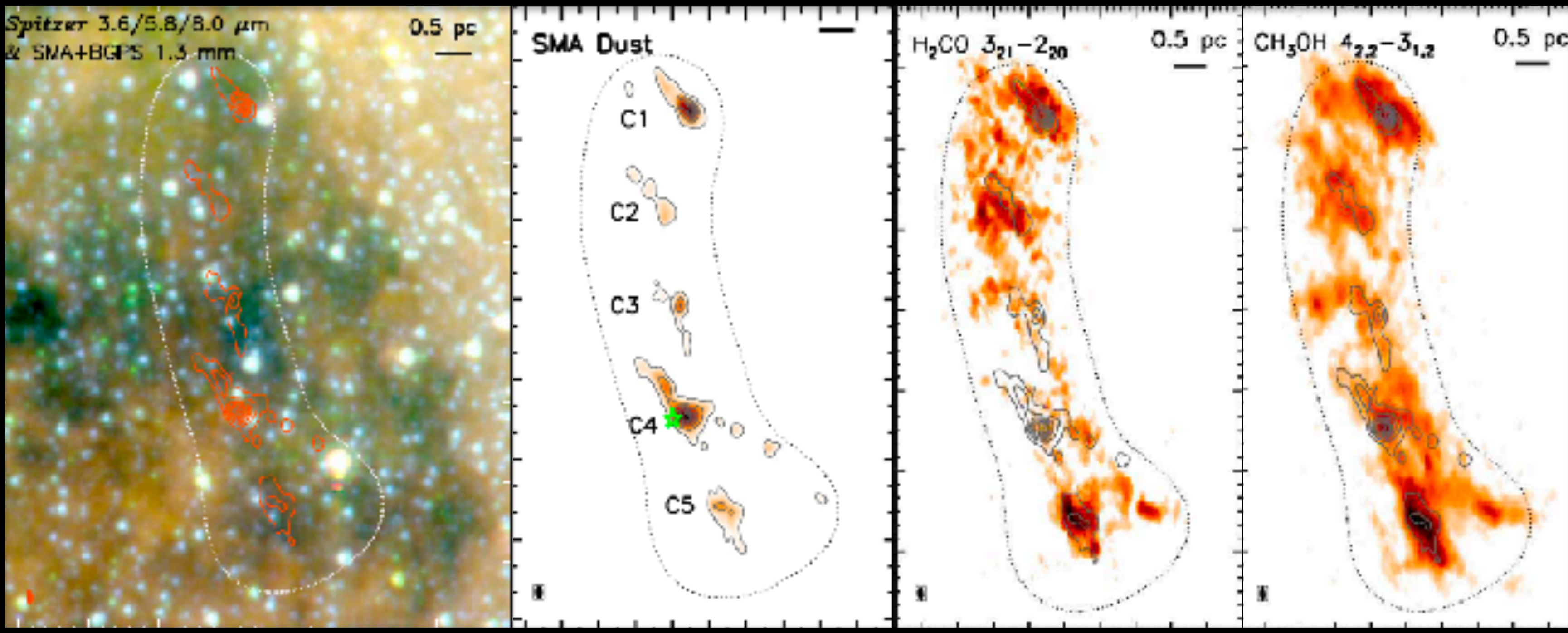


[1]This also heats the gas!

Ginsburg et al. 2016

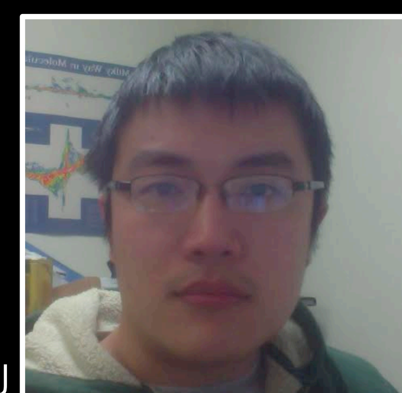
Star Formation in the CMZ

Uncovering Hidden Star Formation



Lu et al 2015, 2017

CMZoom

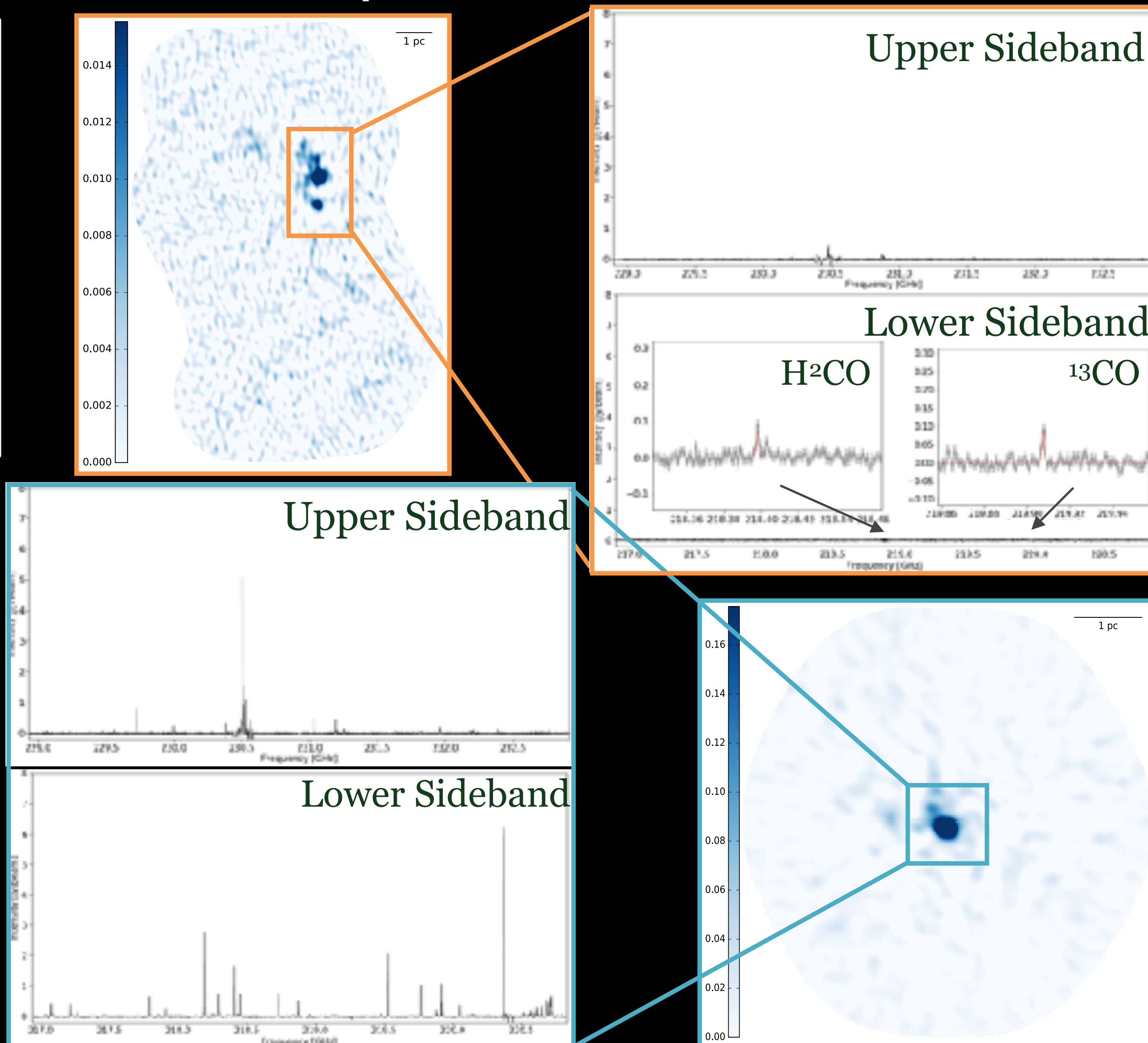


Xing "Walker" Lu
Postdoc at NAOJ

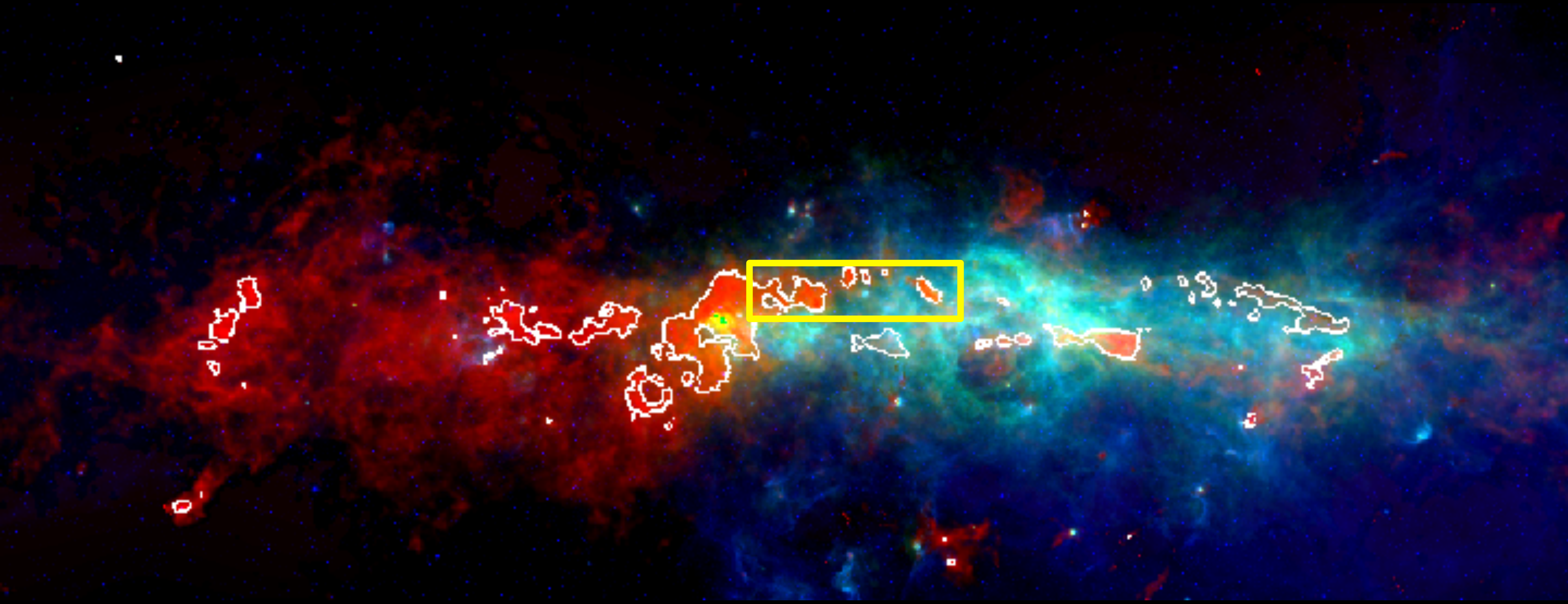
Star Formation in the CMZ

Chemistry in the CMZ

Daniel Callanan
PhD student at
Liverpool/CfA



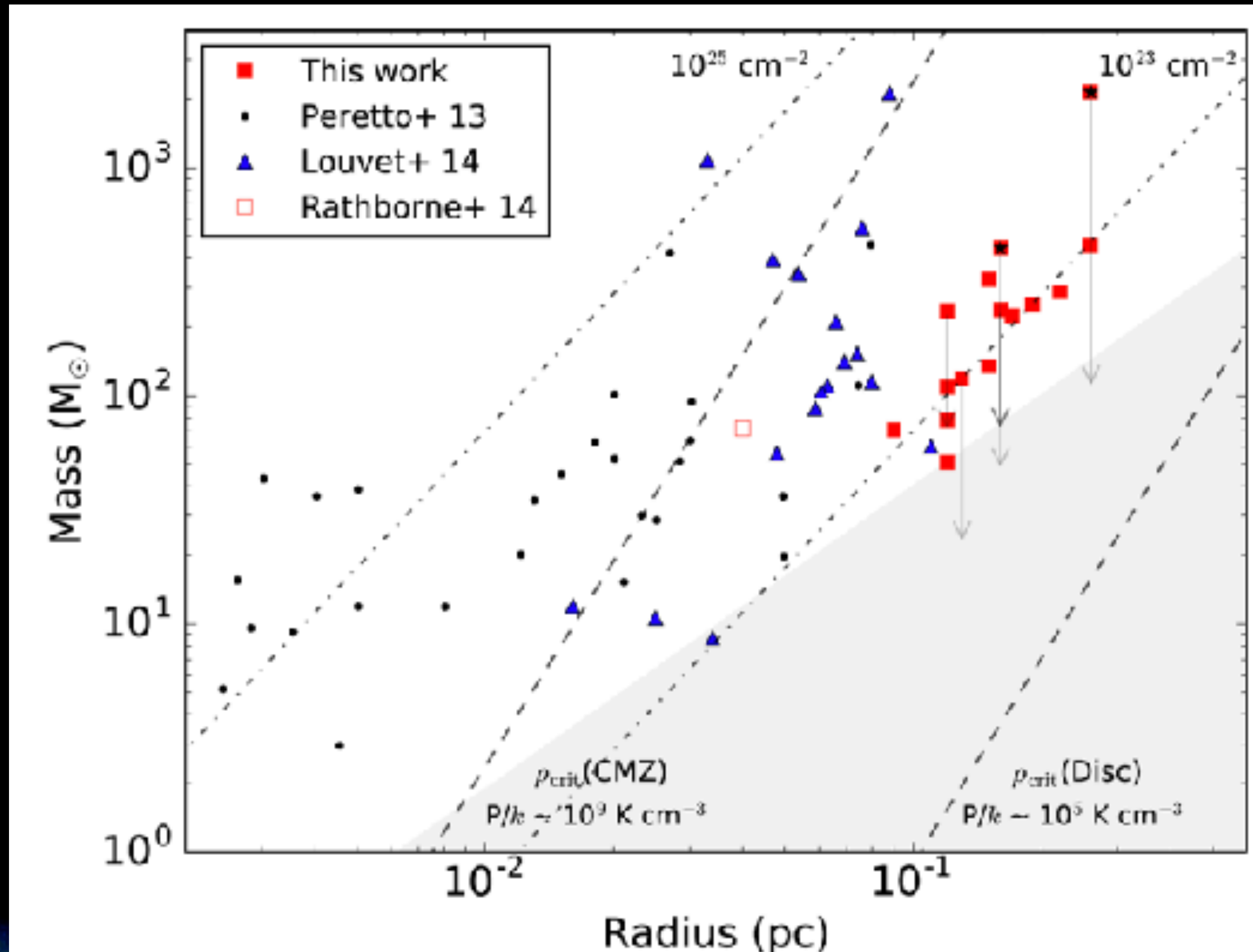
Star Formation in the CMZ



Detailed Study of Core Properties

Dan Walker
postdoc at NAOJ Chile

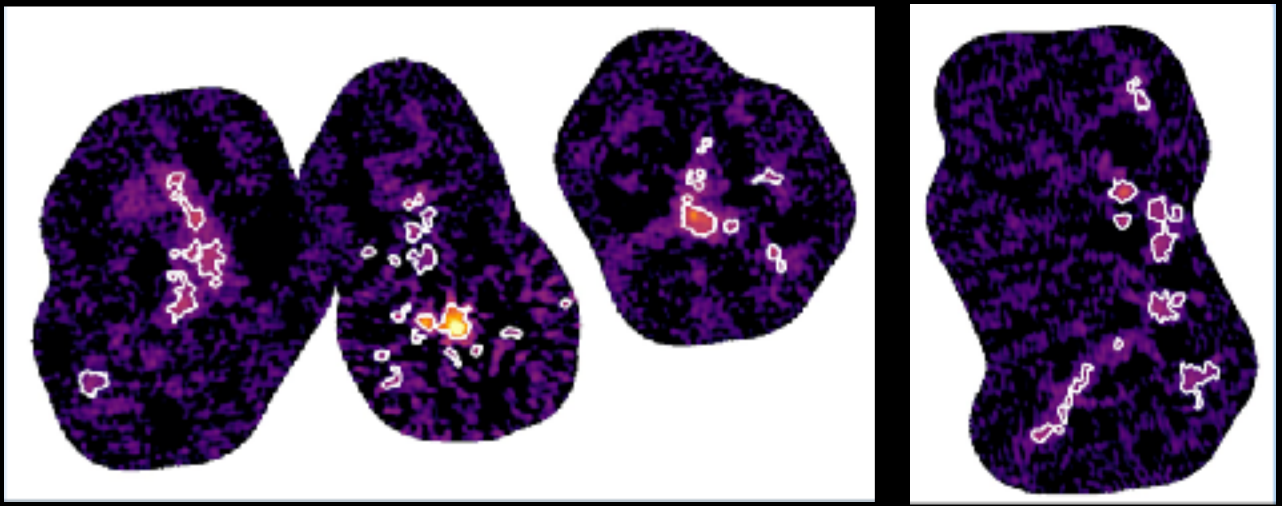
core gas temperatures of about 50-200 K



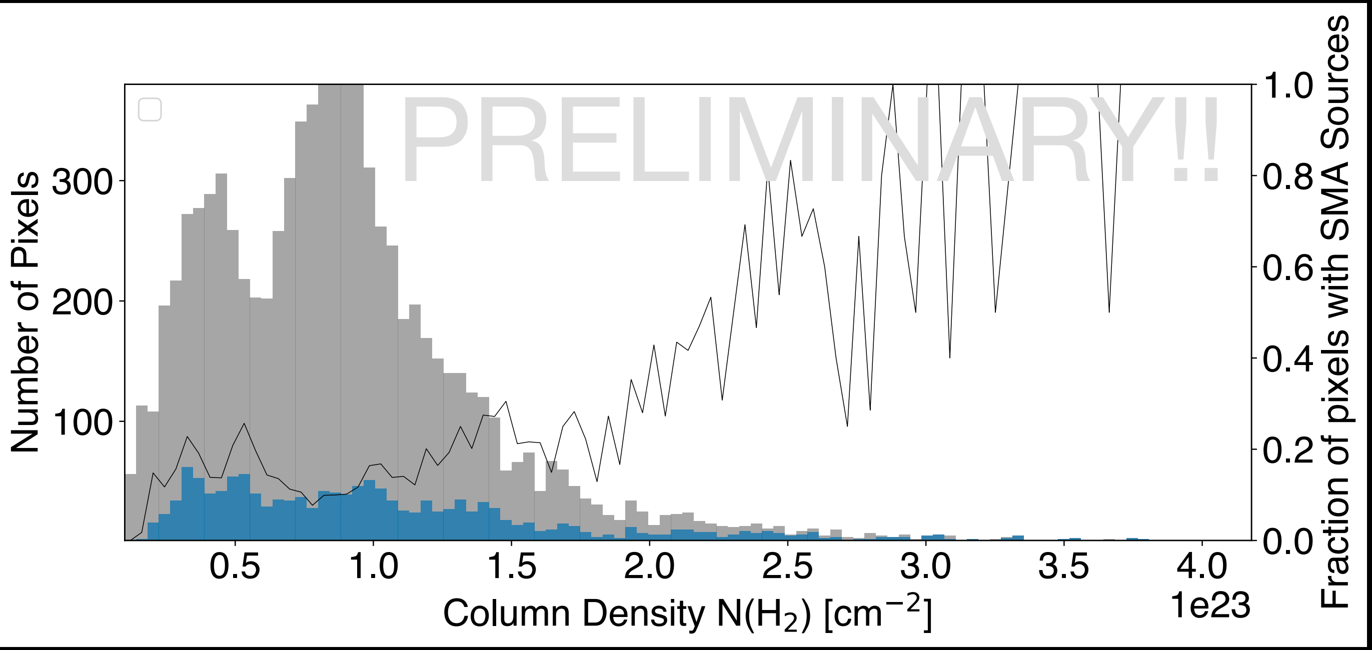
Walker et al. 2018

Catalog and Simulated Observations

Perry Hatchfield
PhD student at
UConn

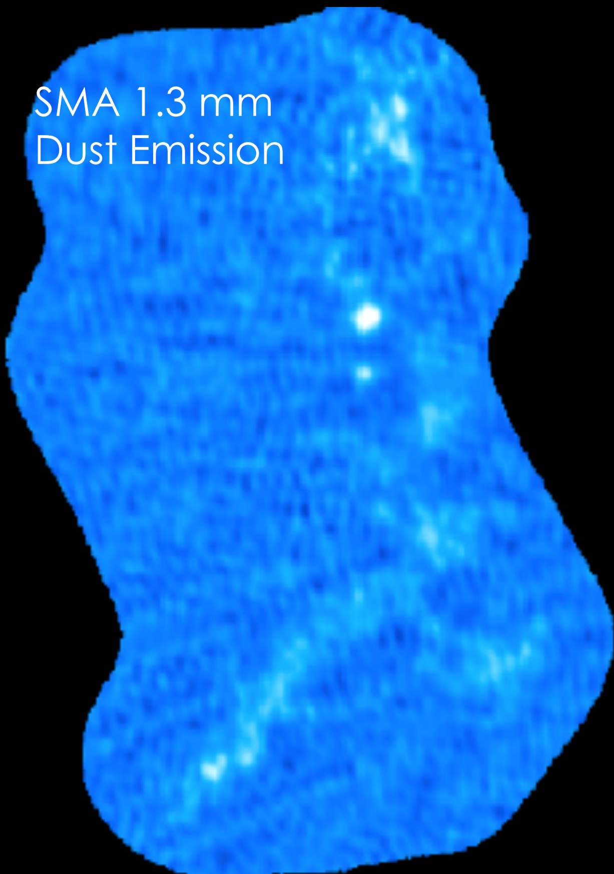


CMZoom Core Catalog: Hatchfield et al. in prep



Star Formation in the CMZ

Why is the SFR so low in the CMZ?



Low-level isolated star formation

Is it star forming?

- ✓ Dense gas
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- ✓ Virial ratio < 2
- ✓ Power-law tail in N-PDF
- ✓ Outflow or localized hot-core chemistry

CMZoom

Battersby et al., in prep.

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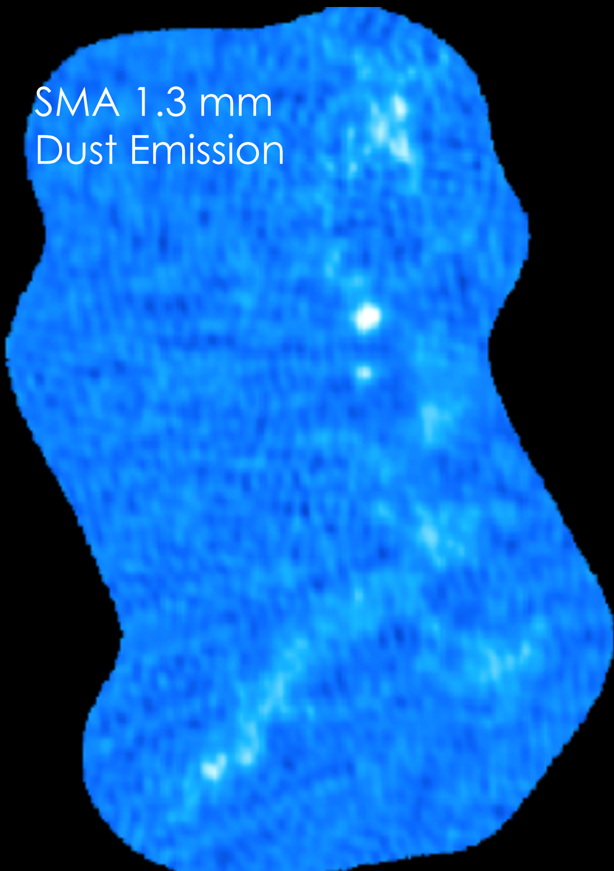

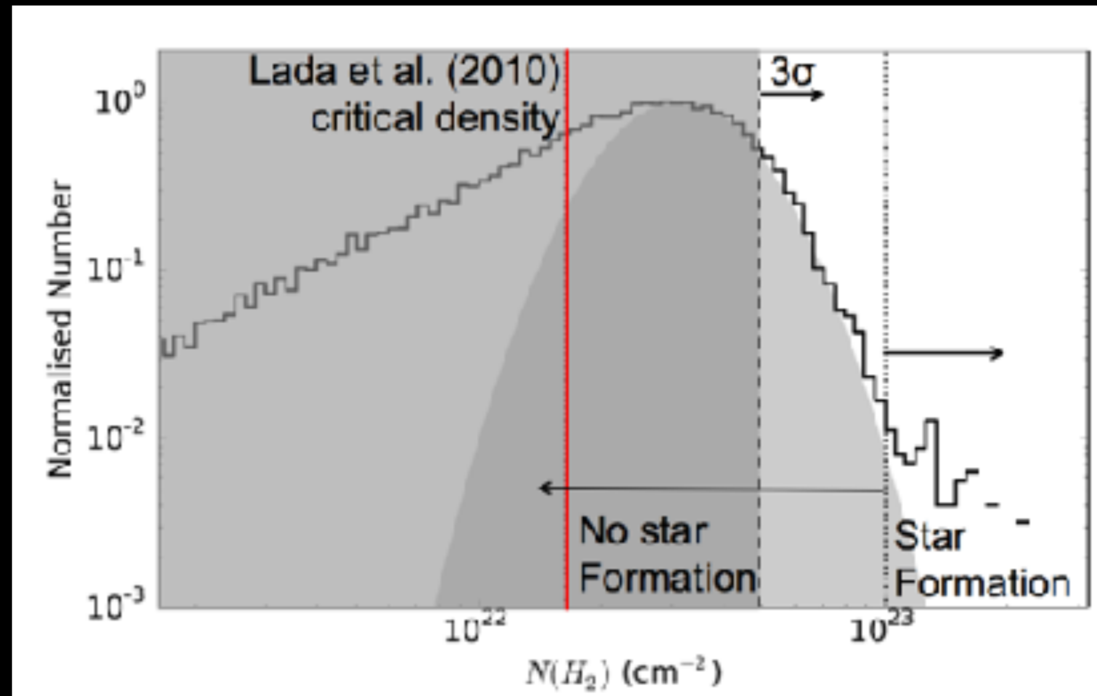

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Battersby et al., in prep.

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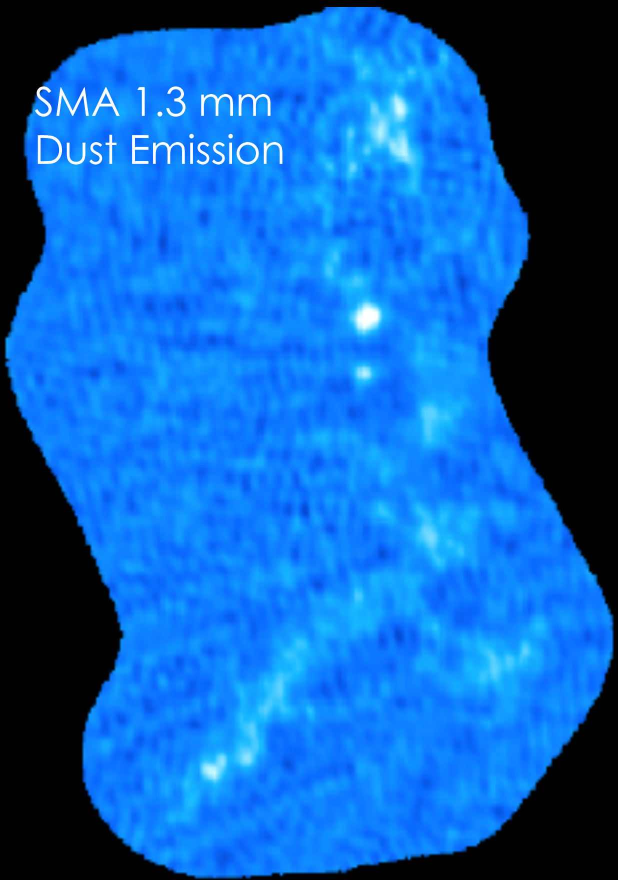

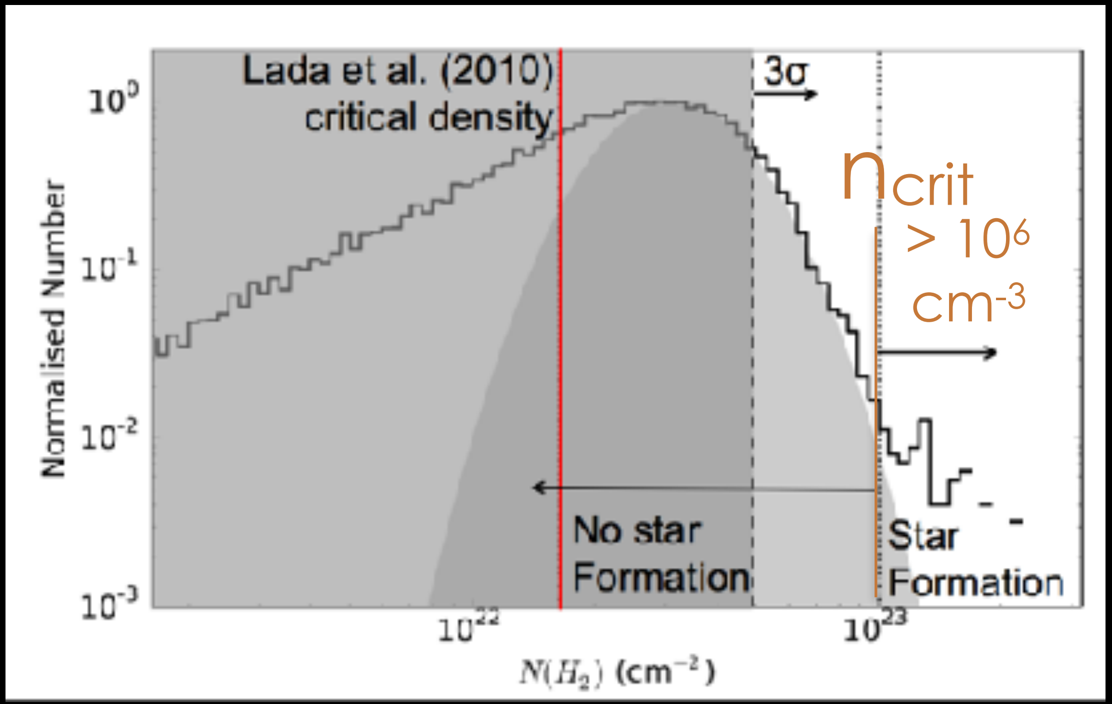

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Battersby et al., in prep.

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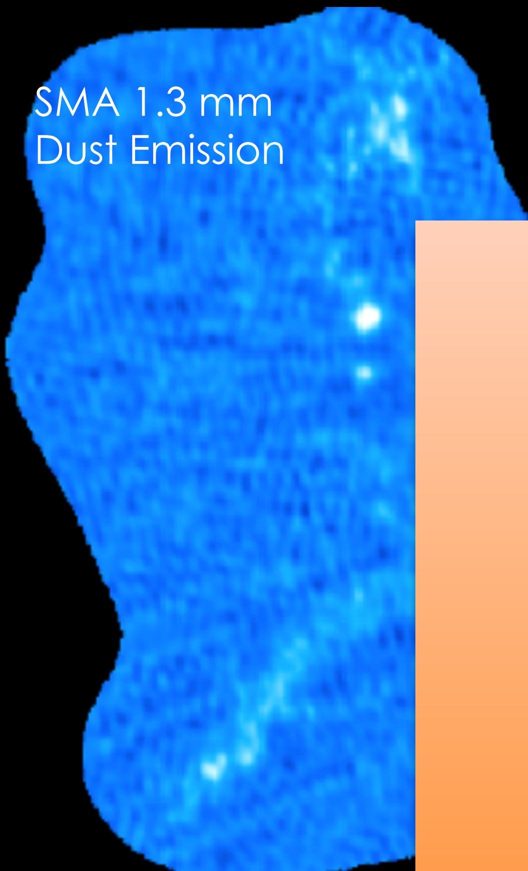
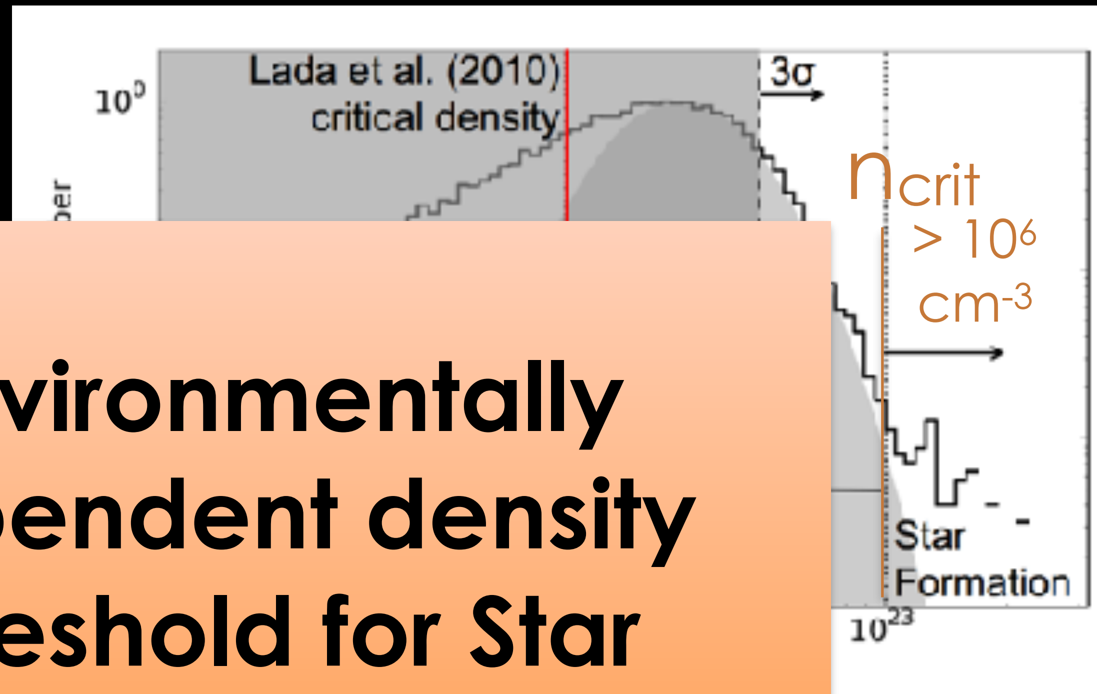




Environmentally Dependent density threshold for Star Formation?

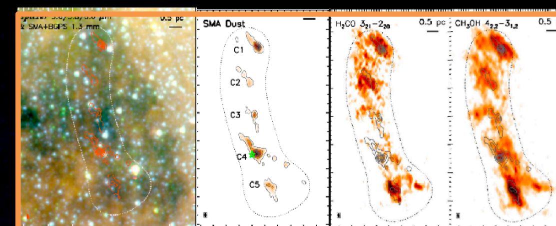
Low-level isolated star formation

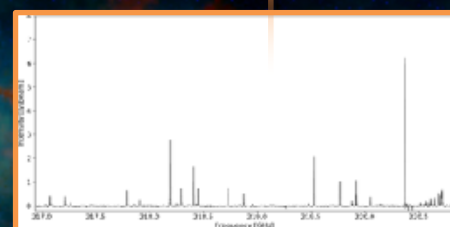
Battersby et al., in prep.

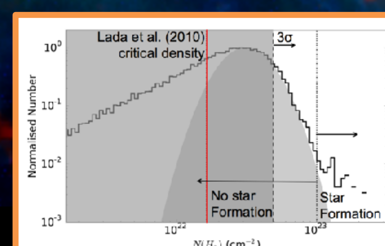
Star Formation in our Extreme Galactic Center: Results from the CMZoom Survey

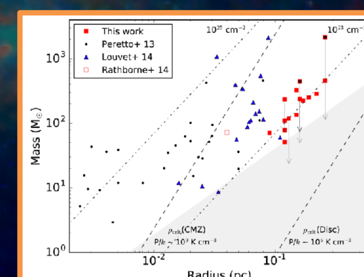
New survey, CMZoom, mapped all the highest column density gas in inner 500 pc and:

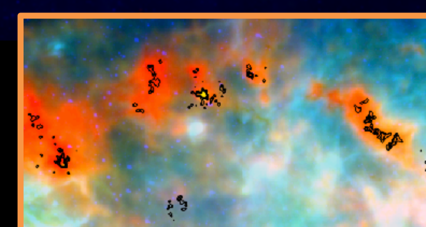
* Uncovered hidden star formation
* CMZ cores demonstrate very different excitation/chemistry
* CMZ cores are on the same mass-radius relation as disk cores
* High levels of turbulence seem capable of inhibiting SF in the CMZ
* Meaning that SFR should depend on environment





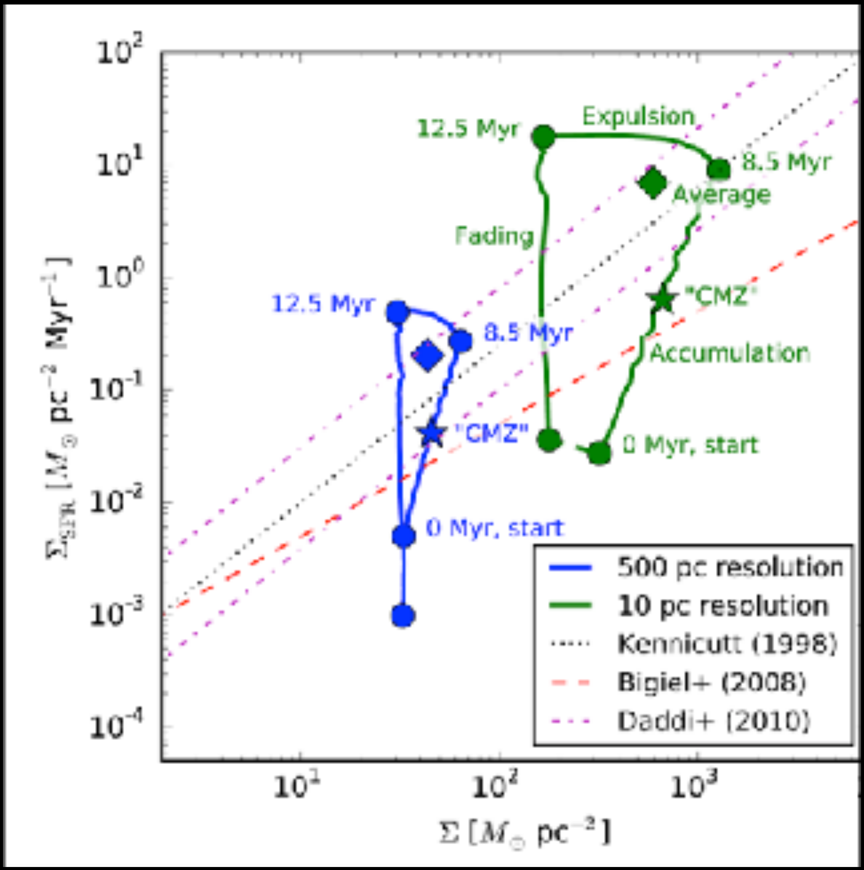






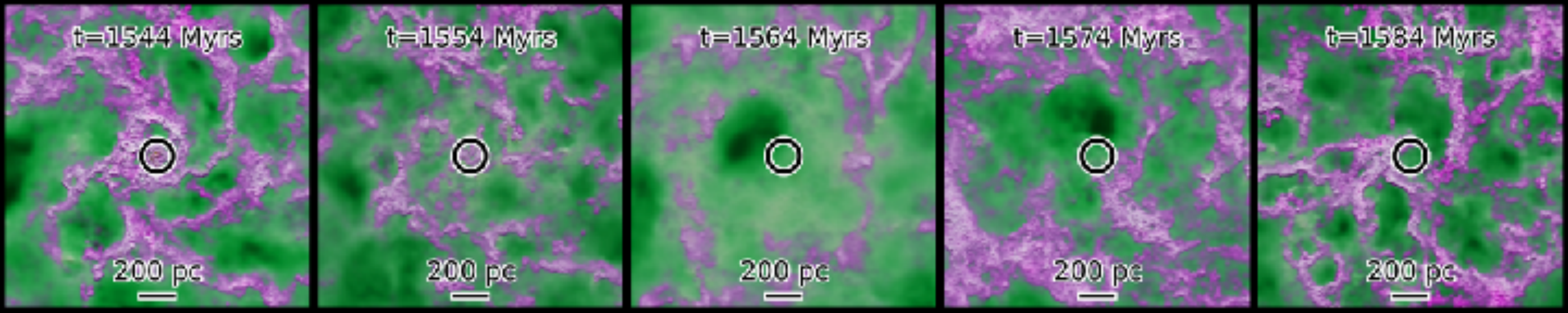
Extra Slides

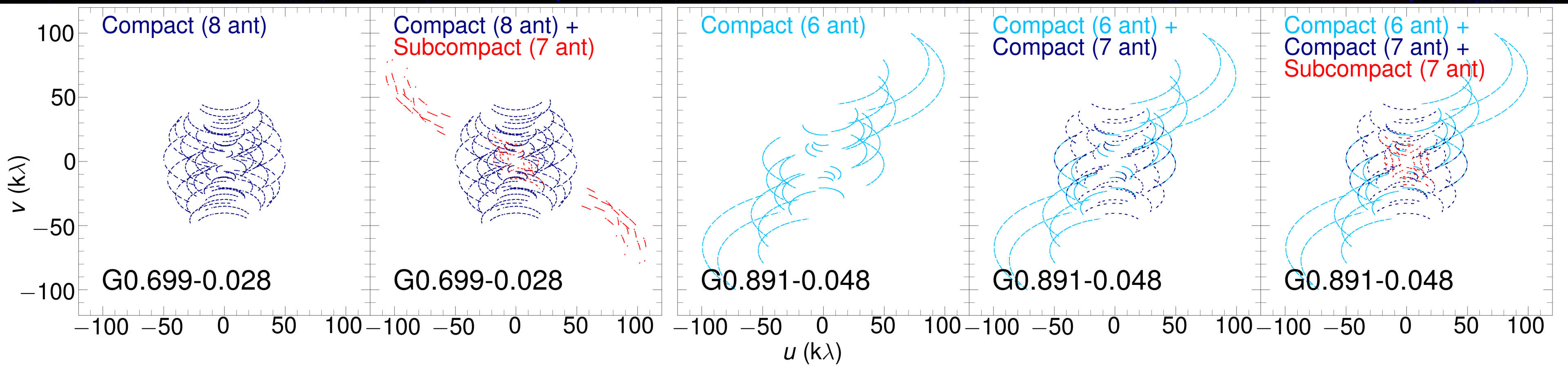
Why is the SFR low in the CMZ?

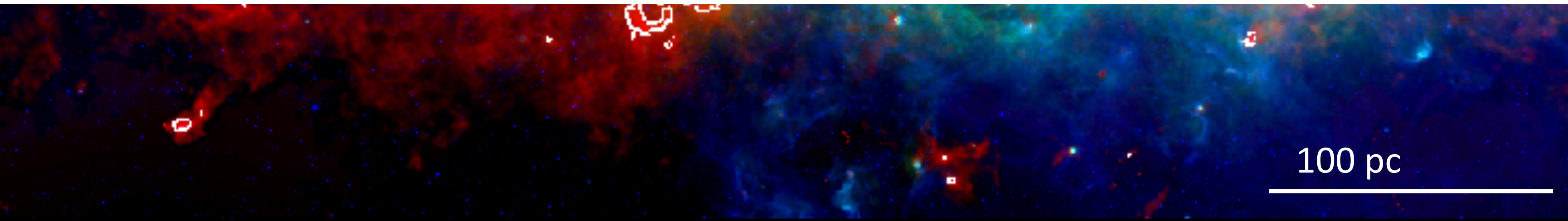


Krumholz & Kruijssen 2015

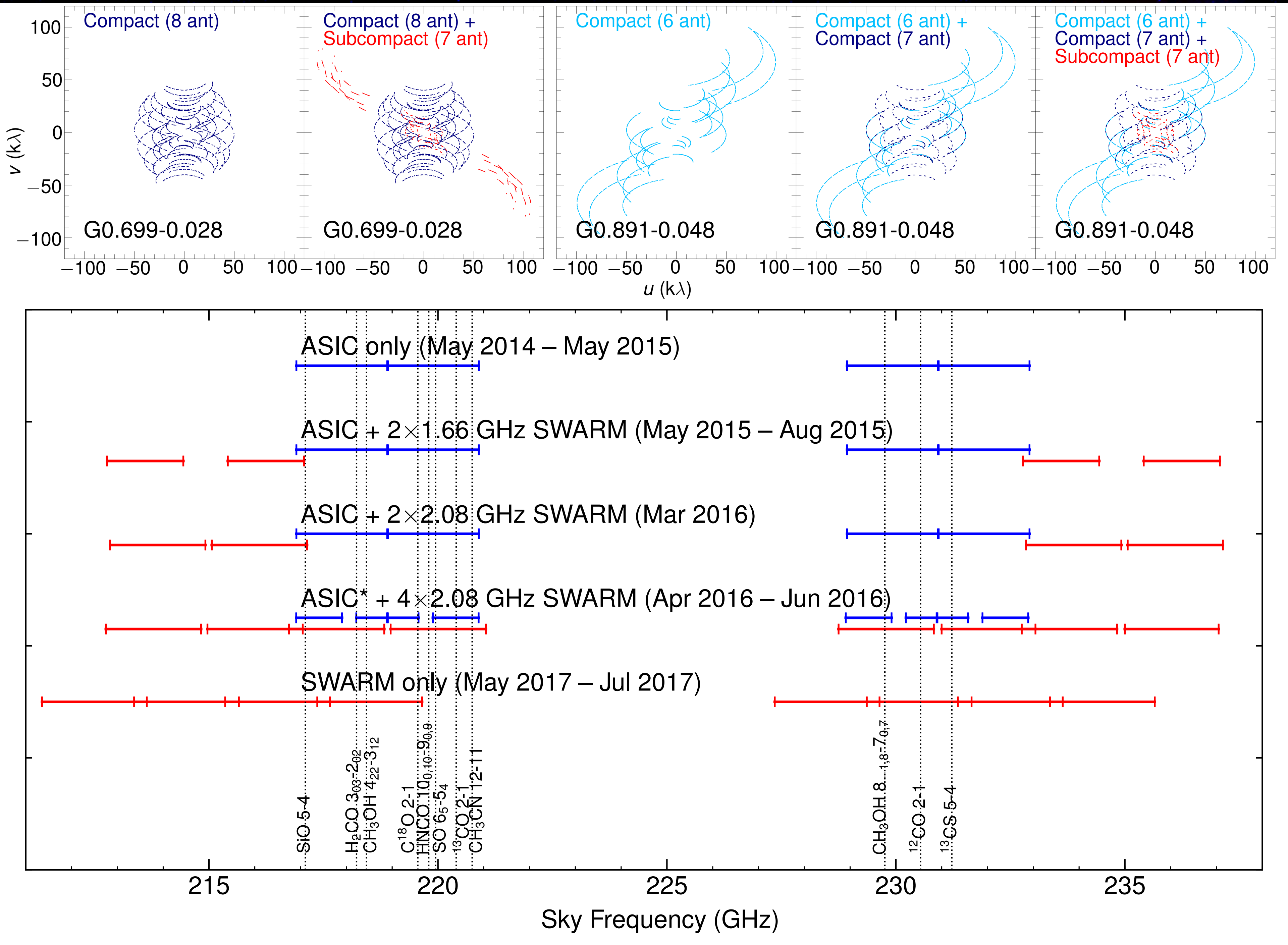
Torrey, Hopkins et al., 2017





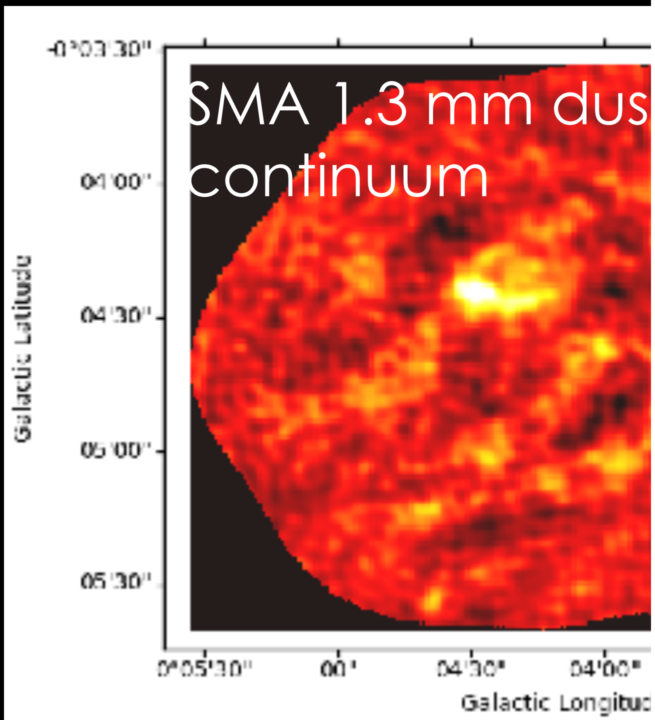


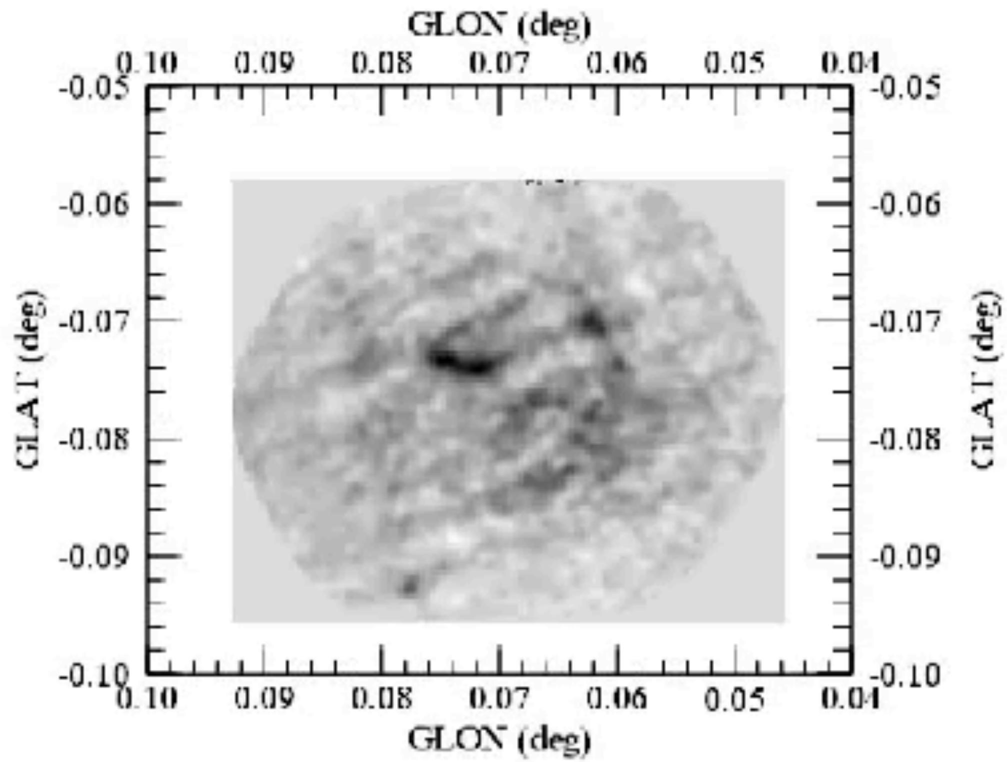
- 230 GHz (1.3 mm)
- 240 arcmin^2 (above N(H_2) = 10^{23} cm^{-2} or 3×10^{22} cm^{-2})
- 4'' (0.2 pc) resolution, $\Delta v \sim 1.1$ km/s
- dust continuum + spectral lines (H_2CO, ^{12}CO, ^{13}CO, $C^{18}O$, SiO, CH_3OH, CH_3CN, etc.): 8+ GHz bandwidth
- 3 mJy RMS continuum, 0.4 K
- 550 hours (50 subcompact, 450 compact/custom) over 4 yrs
- Complement with single-dish (APEX, CSO) observations



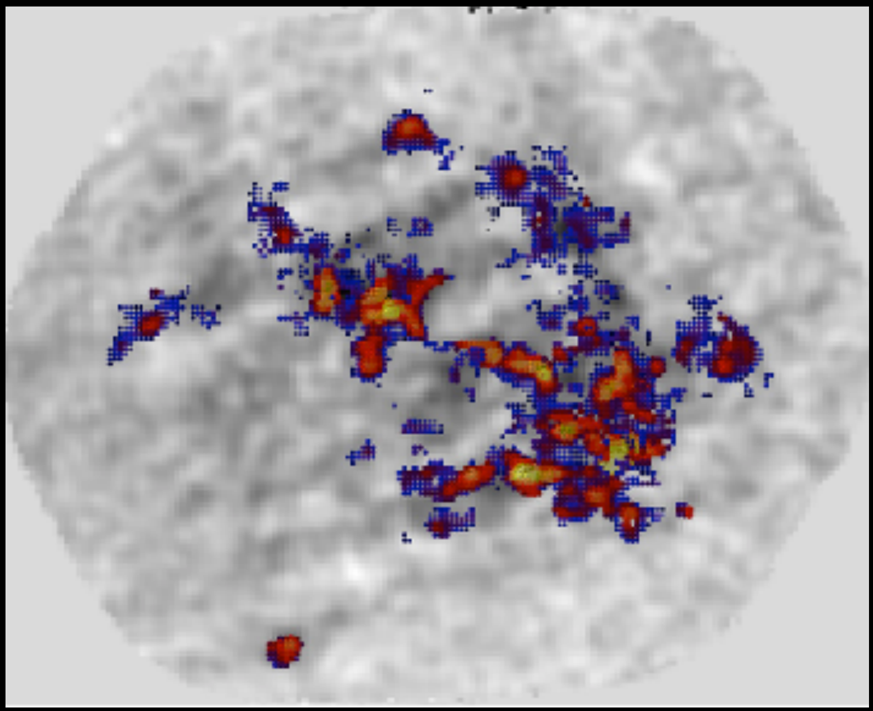
Compact (8 ant)
G0.699-0.028
Compact (8 ant) +
Subcompact (7 ant)
G0.699-0.028
Compact (6 ant)
G0.891-0.048
Compact (6 ant) +
Compact (7 ant)
G0.891-0.048
Compact (6 ant) +
Compact (7 ant) +
Subcompact (7 ant)
G0.891-0.048
v (kλ)
u (kλ)
ASIC only (May 2014 – May 2015)
ASIC + 2×1.66 GHz SWARM (May 2015 – Aug 2015)
ASIC + 2×2.08 GHz SWARM (Mar 2016)
ASIC* + 4×2.08 GHz SWARM (Apr 2016 – Jun 2016)
SWARM only (May 2017 – Jul 2017)
SiO 5-4
H2CO 303-202
CH3OH 422-312
C18O 2-1
HNCO 100,10-90,9
SO 65-54
13CO 2-1
CH3CN 12-11
CH3OH 8−1,8-70,7
12CO 2-1
13CS 5-4
Sky Frequency (GHz)

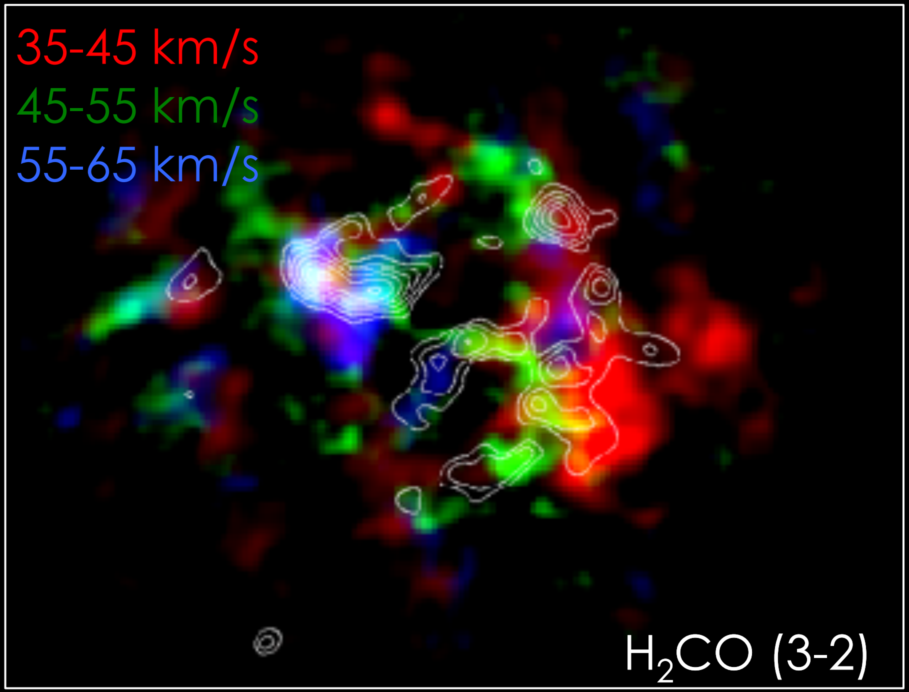
Why is the SFR low in the CMZ?





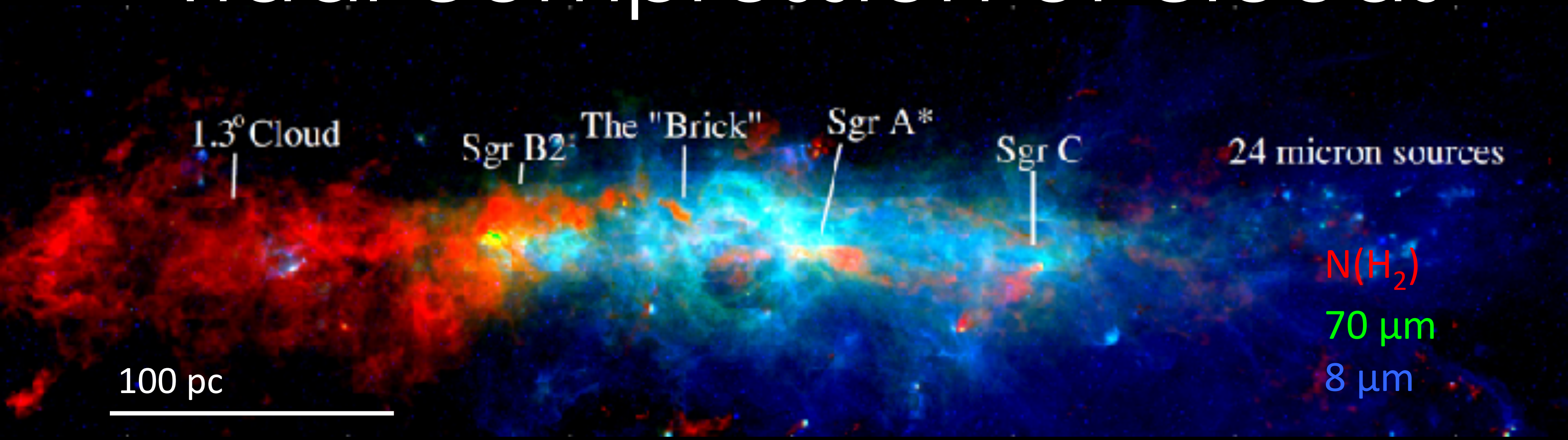
Structure Identification

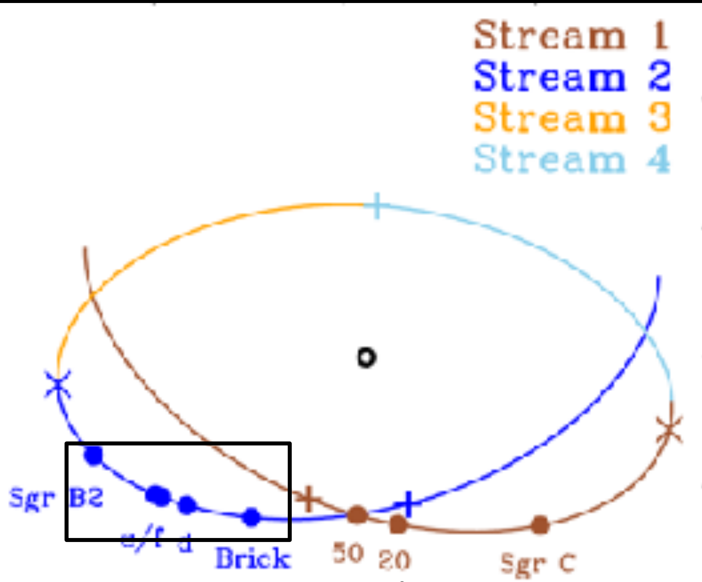




SCOUSE line fitting
Jonny Henshaw
Liverpool

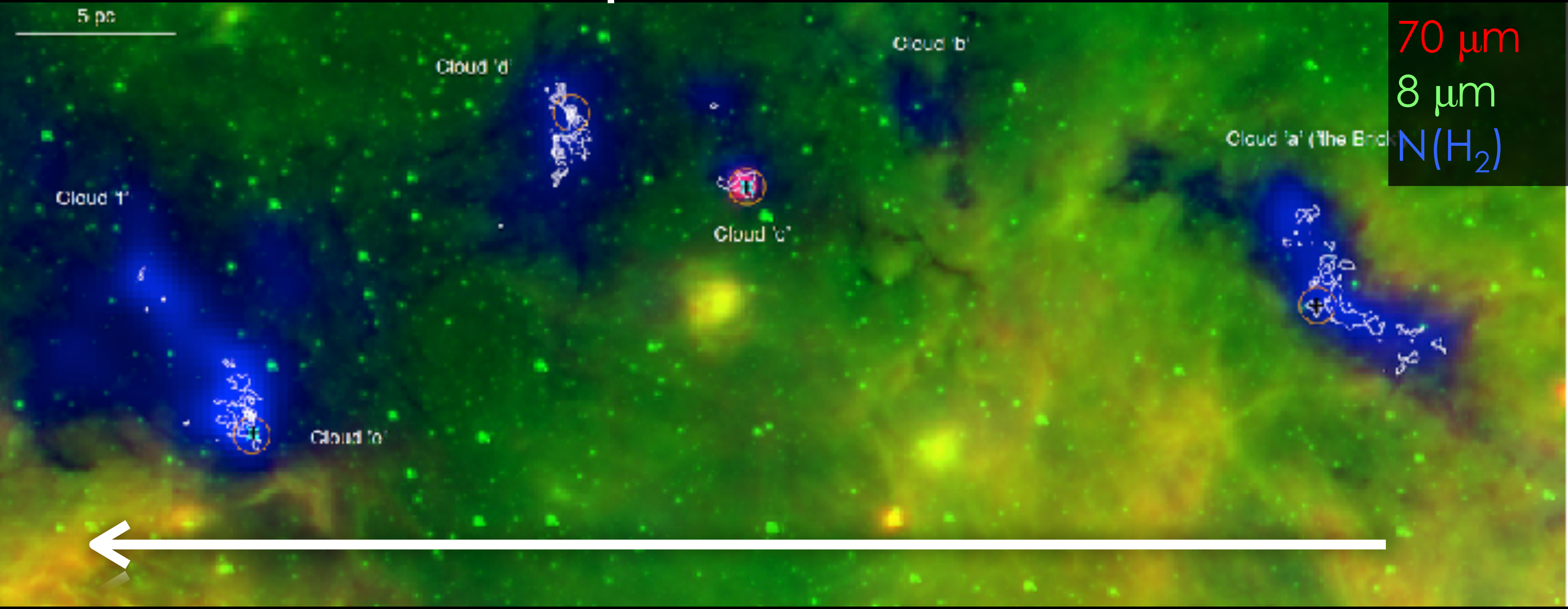
Tidal compression of clouds

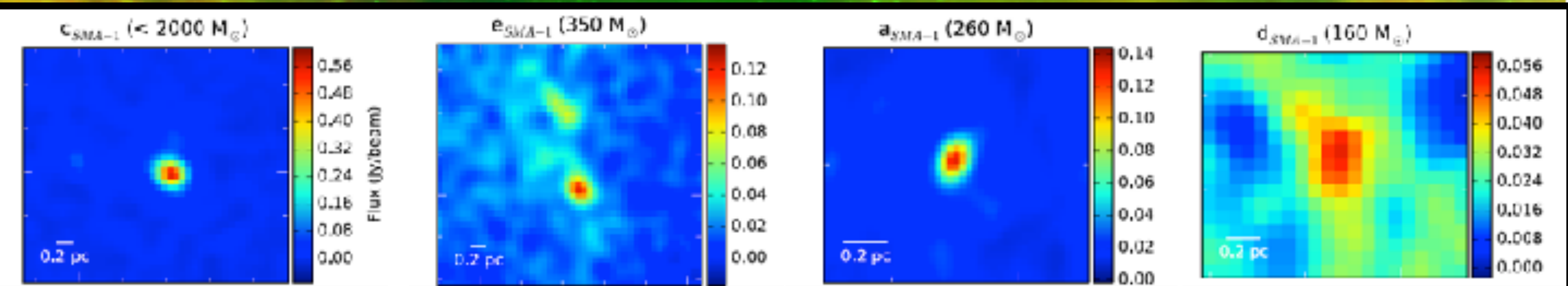




New orbital models (Kruijssen, Dale, & Longmore 2015)

Tidal compression of clouds



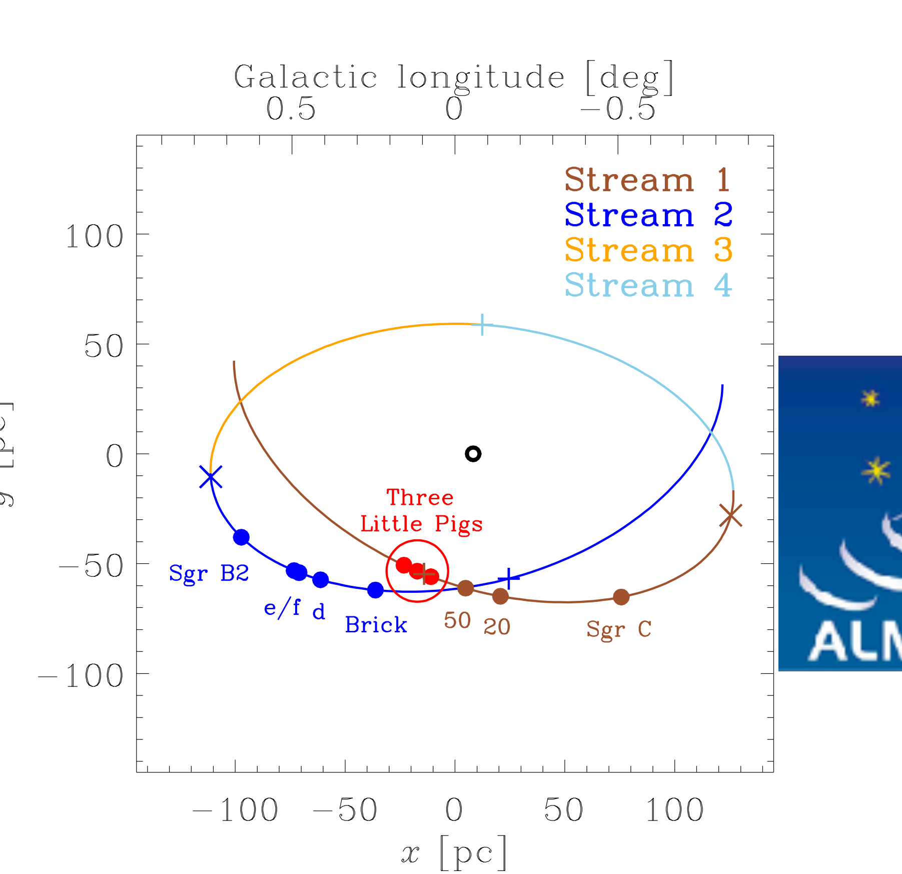


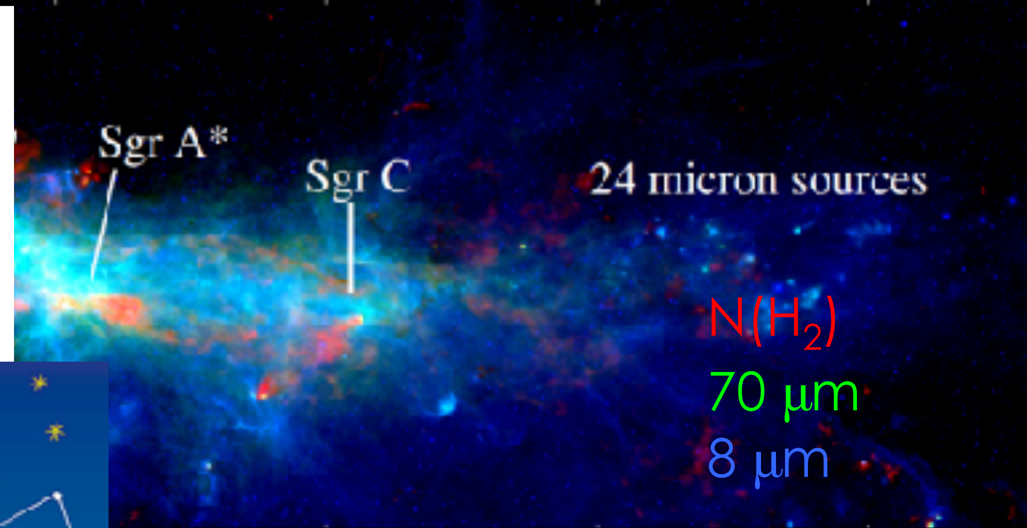
N_{ff} = 1.76
H_2O, CH_3OH, SiO, H_2CO masers + 70 µm

N_{ff} = 1.28
H_2O and CH_3OH masers

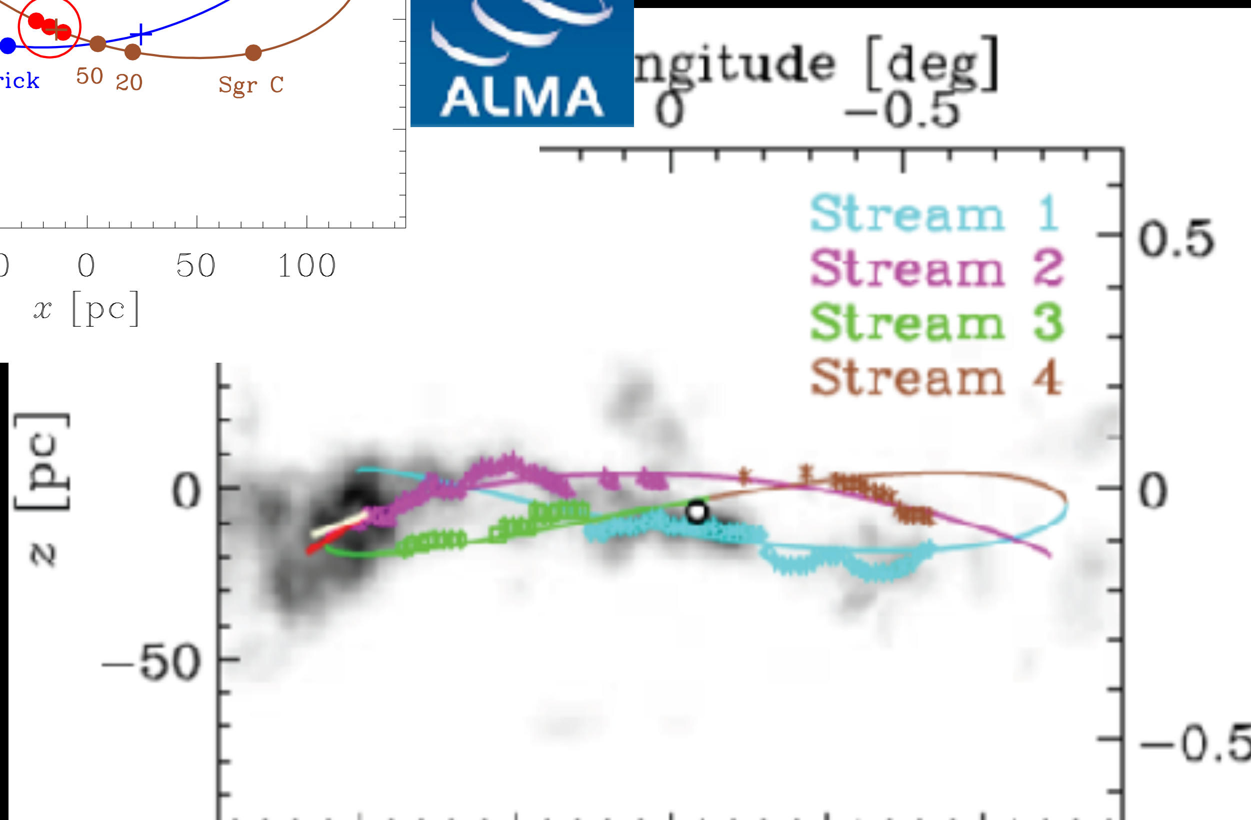
N_{ff} = 0.65
H_2O maser

N_{ff} = 0.61
No SF tracers

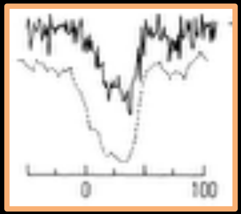




New orbital models (Kruijssen, Dale, & Longmore 2015)

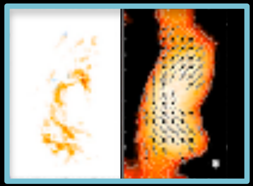


Why is SF low? - Gas properties



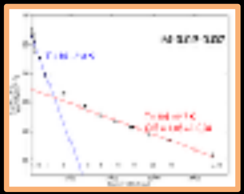
Gas is dense, n> 10^5 cm^{-3}

Guesten et al. 1983, Zylka et al. 1992, Serabyn et al. 1992, Walmsley et al. 1986



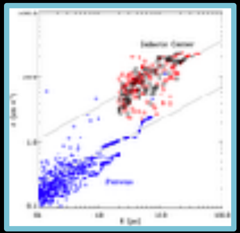
Gas is chemically complex, with ~mG magnetic fields, high ISRF, and high CRIR

complex: e.g. Rathborne et al., 2014, Requena-Torres et al. 2008, **magnetic fields:** e.g. Pillai et al. 2015, Yusef-Zadeh & Morris 1987, **high ISFR and CRIR:** e.g. Clarke et al. 2013, Goto et al. 2013, etc.



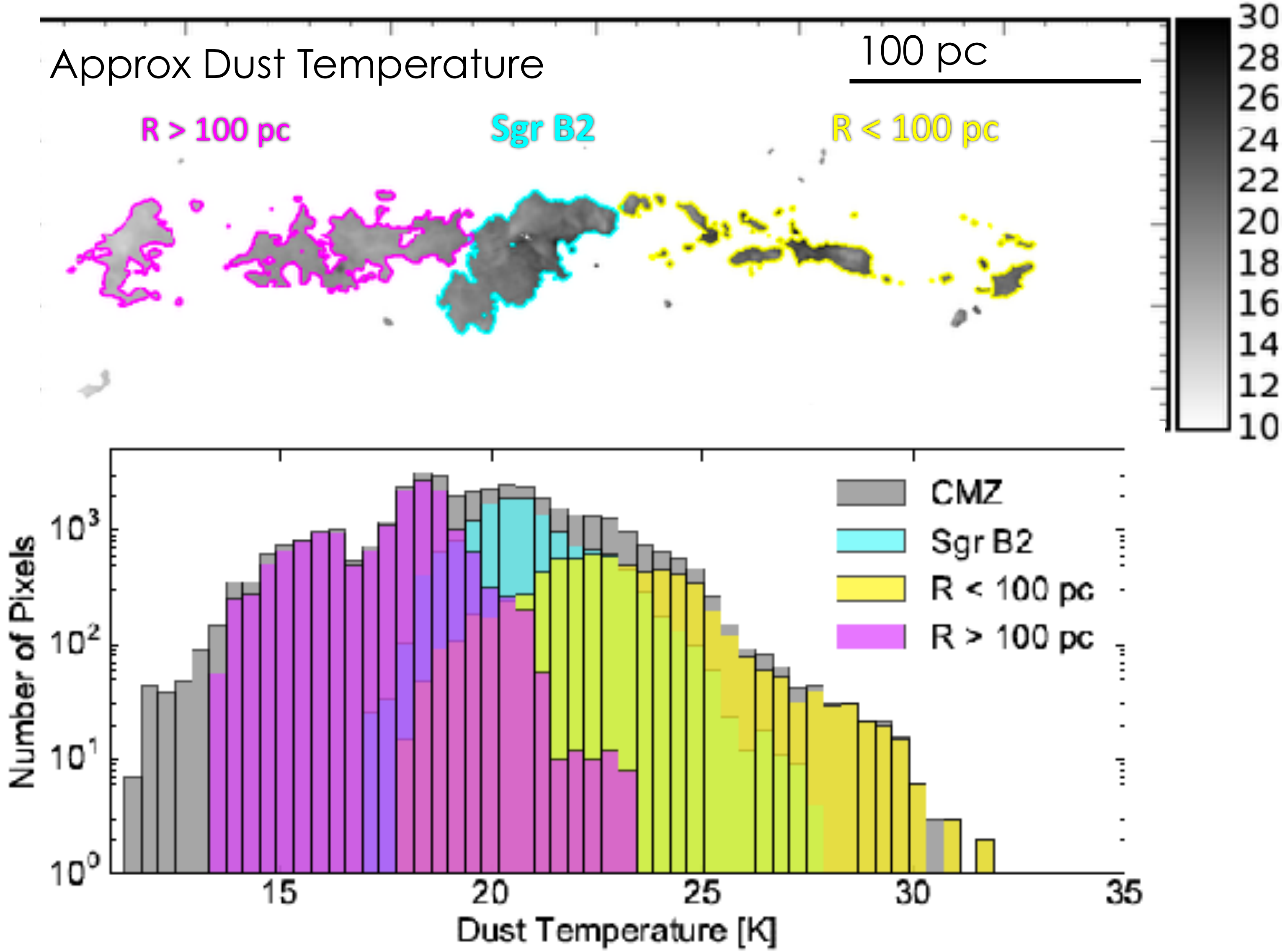
The dense gas is hot (>65 K), and 10% is 400 K

Ginsburg et al. 2016, Mills et al. 2013, Immer et al. 2016, Ott et al. 2014, Krieger et al. in prep.

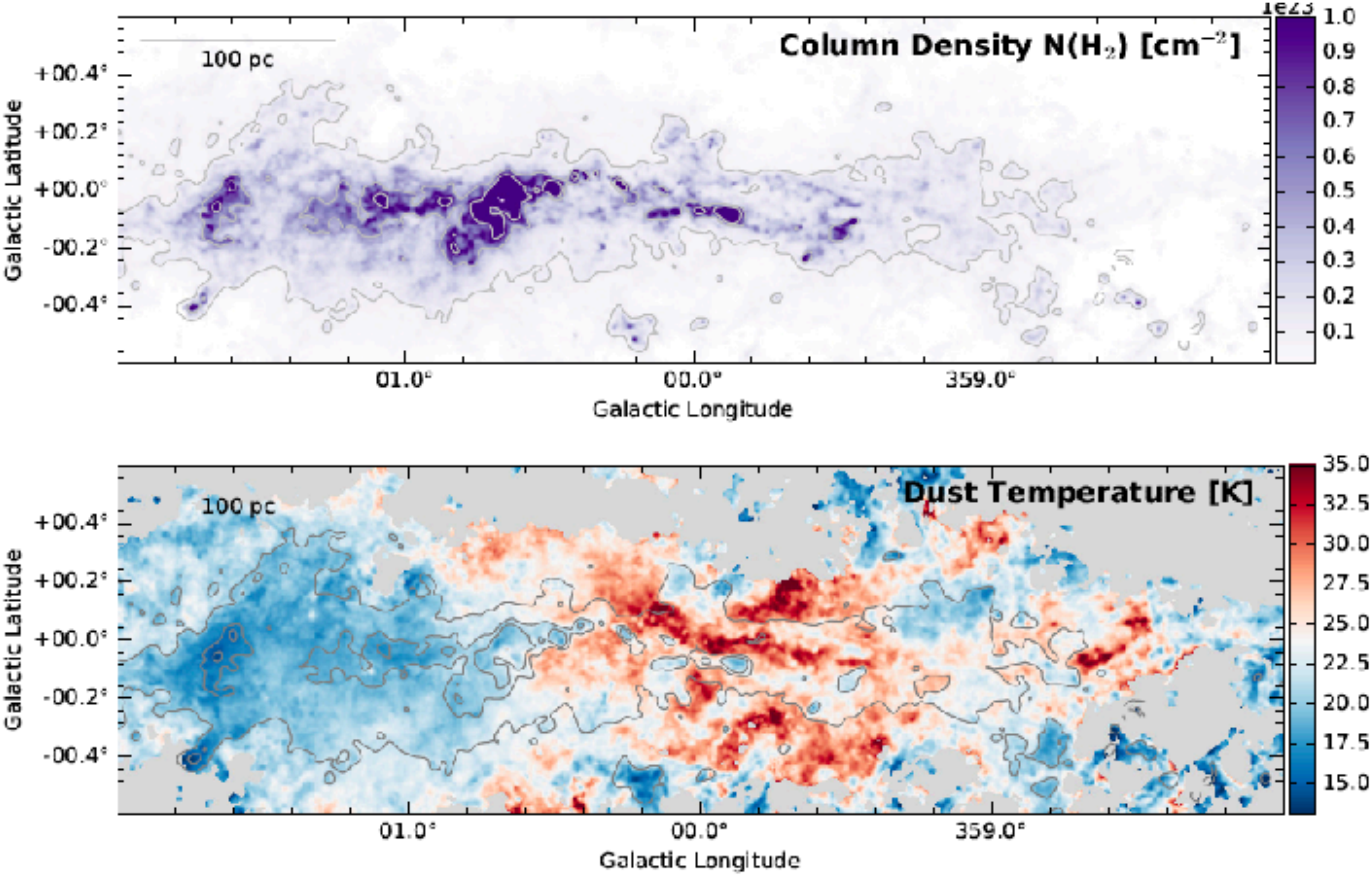


The gas is turbulent (Δv ~10 km/s, $\mathcal{M}$~10-40)

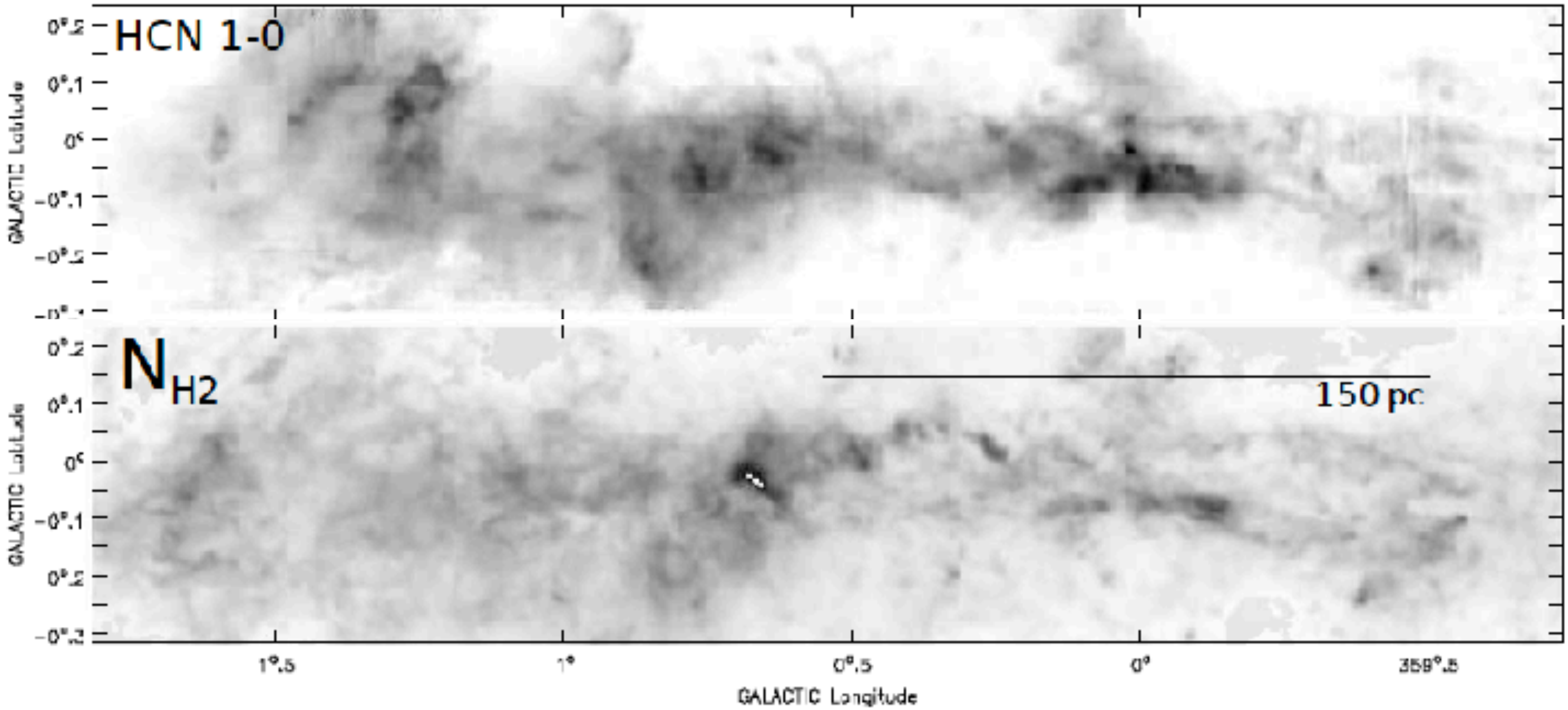
Shetty et al. 2012, Rathborne et al. 2015, Kauffmann et al. 2017 Ginsburg et al. 2016, Mills et al. 2013, Immer et al. 2016, etc.



Battersby, Longmore, Bally, Barnes, Ginsburg, & Mills in prep.



Battersby, Longmore, Bally, Barnes, Ginsburg, & Mills in prep.



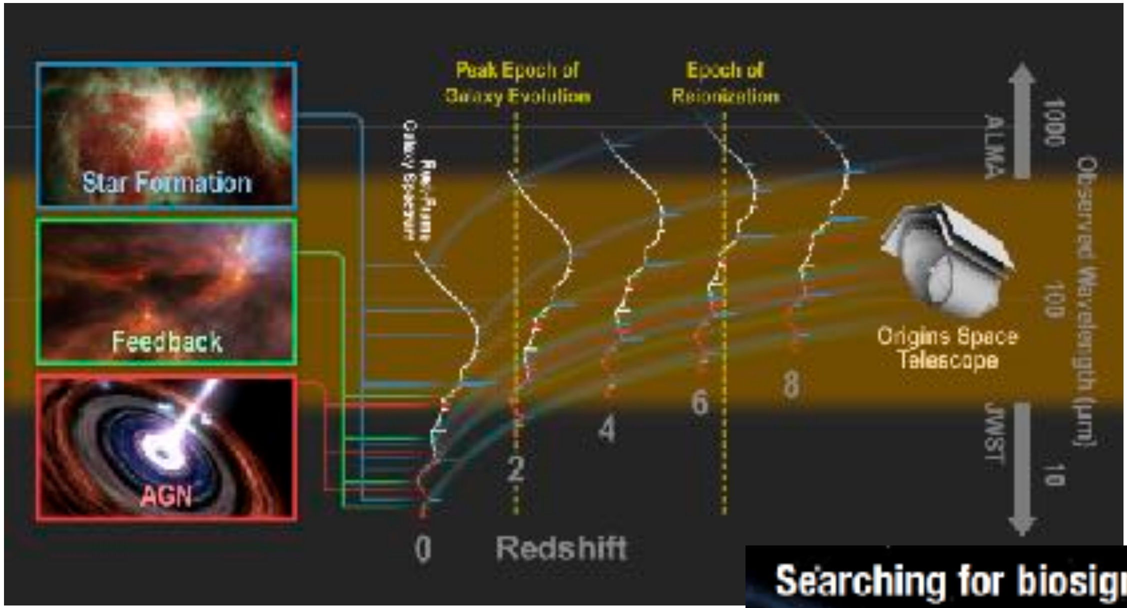
- HCN is well-correlated with dense gas overall in the CMZ - variations would only yield a 10% error in the dense gas mass
  - However, there is a lot of scatter (0.75 dex)
  - Some clouds are under-bright or over-bright by factors of 2-3
  - This is bad if you are looking at an AGN or shock-dominated region of a galaxy
- A lot of the HCN comes from more diffuse gas
- HNCO is better correlated with dense gas

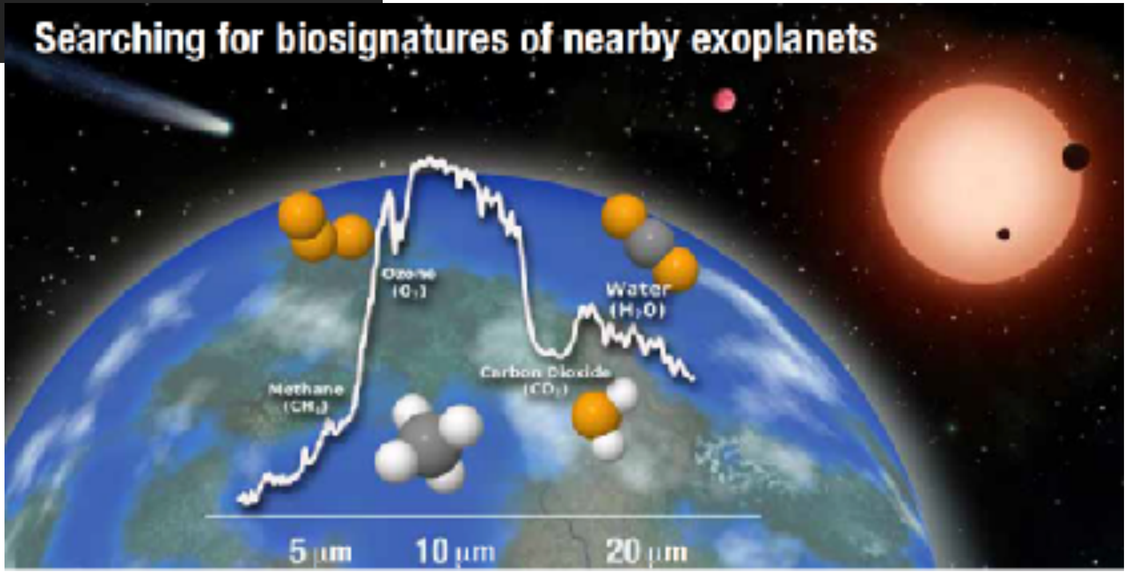
Mills & Battersby 2017

Next Generation CMZoomers

The Milky Way Laboratory at UConn
Battersby Research Group Fall 2017

ORIGINS
Space Telescope





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